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VLT AND ACS OBSERVATIONS OF RDCS J1252.9-2927: DYNAMICAL STRUCTURE AND GALAXY POPULATIONS IN A MASSIVE CLUSTER AT Z=1.237*

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ABSTRACT

We present results from an extensive spectroscopic survey, carried out with FORS on the ESO Very Large Telescope, and from an extensive multi-wavelength imaging data set from the Advanced Camera for Surveys and ground based facilities of the cluster of galaxies RDCS J1252.9-2927. We have spectroscopically confirmed 38 cluster members in the redshift range $1.22 < z < 1.25$. The distribution in velocity of these spectroscopic members yields a cluster median redshift of $z = 1.237$ and a rest-frame velocity dispersion of 747^{+74}_{-84} km s^{-1}. Star-forming members are observed to mainly populate the outskirts of the cluster while passive galaxies dominate the central cluster region. Using the 38 confirmed redshifts, we were able to resolve, for the first time at $z > 1$, kinematic structure. The velocity distribution, which is not Gaussian at the 95% confidence level, is consistent with two groups that are also responsible for the projected elongation of the cluster in the East-West direction. The groups are composed of 26 and 12 galaxies and have velocity dispersions of 486^{+47}_{-85} km s^{-1} and 426^{+57}_{-105} km s^{-1}, respectively. The elongation is also seen in the intracluster gas (from X-ray observations) and the dark matter distribution (from a weak lensing analysis). This leads us to conclude that RDCS J1252.9-2927 has not yet reached a final virial state. We extend the analysis of the color-magnitude diagram of spectroscopic members to more than 1 Mpc from the cluster center. The scatter and slope of non-[OII]-emitting cluster members in the near-IR red sequence is similar to that seen in clusters at lower redshift. Furthermore, none of the galaxies with luminosities greater than $\sim K^*_s + 1.5$ show any [OII] emission feature, indicating that these more luminous, redder galaxies have stopped forming stars earlier than the fainter, bluer galaxies. Our observations provide detailed dynamical and spectrophotometric information of galaxies in this exceptional high-redshift cluster, delivering an in-depth view of structure formation at this epoch only 5 Gyr after the Big Bang. In addition, we present the spectra of a few X-ray point sources in the cluster field of view.

Subject headings: galaxy clusters: general — galaxy clusters: individual(RDCS J1252.9-2927)

1. INTRODUCTION

Clusters of galaxies are tracers of the peaks of matter density in the universe. Their study over a wide range in redshift provides an insight into the process of mass assembly of structures through cosmic history, from galactic to Mpc scales. They are suitable laboratories in which to study galaxy populations, providing clues to better understand the effects of the local environment on galaxy properties and galaxy evolution. Cluster X-ray luminosity and temperature indicate the existence of massive systems ($M > 10^{14} M_\odot$) already when the universe was half its present age (Jeltema et al. 2001; Maughan et al. 2003, 2004; Rosati et al. 2004; Mullis et al. 2005). While some of these clusters have a relaxed X-ray morphology by $z \sim 0.8$ (Maughan et al. 2004), others show clear signatures of being still in a formation stage due to their filamentary morphology and the presence of substructure (Gioia et al. 1999, 2004; Demarco et al. 2005; Girardi et

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al. 2005; Tanaka et al. 2006).

The dense cluster environment has a profound effect on the properties of the cluster galaxies. Most notably, early-type galaxies are more prevalent in clusters than in low-density, field environments (Dressler 1980), and this morphology-density relation has existed since $z \sim 1$ (Dressler et al. 1997; Smith et al. 2005; Postman et al. 2005). The existence of the morphology-density relation suggests that the properties of the stellar populations of galaxies also depend on their environment, which is corroborated by the relation between star-formation history, color, and environment (Balogh et al. 1998; Helsdon & Ponman 2003; Kauffmann et al. 2004; Wake et al. 2005; Postman et al. 2005; Homeier et al. 2006a; Thomas & Katgert 2006).

Measurements of the evolution of the mass-to-light ratio of massive cluster early-type galaxies have demonstrated that their stellar populations were largely formed at $z > 2$ (e.g., van Dokkum & Stanford 2003; Holden et al. 2005). The slow evolution of the color-magnitude relation (e.g., Blakeslee et al. 2003b; Holden et al. 2004; Blakeslee et al. 2006; Mei et al. 2006a,b; Homeier et al. 2006b) shows that this is the case for early-type galaxies well below the characteristic luminosity although there seems to be a lack of faint, red galaxies at redshifts higher than $z \sim 0.5$ (Tanaka et al. 2005). Remarkably, massive field early-type galaxies have been shown to be not much younger, in terms of their stellar populations, than cluster early-type galaxies with the same mass (van der Wel et al. 2005; Treu et al. 2005; van Dokkum & van der Marel 2006). This raises the question whether galaxy mass or environment is the most important driver of galaxy evolution.

Most of our knowledge of cluster evolution comes from observations at redshift lower than unity and a few sparse data sets at z = 1-1.3. Observing galaxy clusters at $z > 1$ is difficult and only a small number of galaxy clusters at such a high redshift have been confirmed to date (Stanford et al. 1997; Rosati et al. 1999; Stanford et al. 2002; Rosati et al. 2004; Hashimoto et al. 2004; Mullis et al. 2005; Stanford et al. 2005, 2006; Brodwin et al. 2006; Elston et al. 2006). With the primordial activity of cluster and massive galaxy formation happening at $z \gtrsim 1$, the observations of systems at those redshifts offer the possibility of learning more about their physical properties at an epoch when baryons are first being assembled into gravitationally bound systems within massive dark matter halos. One of those high redshift clusters, RDCS J1252.9-2927 (Rosati et al. 2004), is a massive X-ray luminous system at $z = 1.24$ selected from the ROSAT Deep Cluster Survey (RDCS; Rosati et al. 1998), which has been the center of an intensive multi-wavelength campaign in the past 5 years.

RDCS J1252.9-2927 ($\alpha_{J2000} = 12^h52^m48^s$, $\delta_{J2000} = -29^o27'00''$) was discovered in the 15.7 ksec ROSAT PSPC field with ID WP300093 at an off-axis angle of 13.′9 with 31 net counts, corresponding to a flux of $(2.5 \pm 0.9) \times 10^{-14}$ erg cm^{-2} s^{-1} in the 0.5-2 keV band. I-band imaging (30 minutes exposure), obtained with the Prime Focus camera at the CTIO 4-m telescope in February 1997, revealed a faint ($I \simeq 21.7$) galaxy pair very close to the X-ray centroid position. In November 1998, J- and K_s-band imaging obtained with SofI (Moorwood et al. 1998) on the ESO New Technology Telescope (NTT) showed a clear overdensity of red galaxies with $J - K_s \simeq 1.9$, typical of early-type galaxies at $z > 1$ (Lidman et al. 2004).

X-ray imaging with Chandra and XMM-Newton has been obtained allowing a complete modeling of the X-ray surface brightness of the cluster together with an unprecedented accuracy in the estimation of its temperature, metallicity and dynamical mass (Rosati et al. 2004). RDCS J1252.9-2927 is a massive structure with a luminosity of $L_x(bol) = (6.6 \pm 1.1) \times 10^{44}$ erg s^{-1} (H_0 = 70 km s^{-1} Mpc^{-1}, Ω_m=0.3, and Ω_λ=0.7) measured within an aperture of 60″(500 kpc) radius and a total mass of $M = (1.9 \pm 0.3) \times 10^{14} M_\odot$ within a R_{500} radius[19] of 536 ± 40 kpc (Rosati et al. 2004). Its intracluster medium (ICM) is characterized by a temperature $T_x = 6.0^{+0.7}_{-0.5}$ keV and a metallicity $Z = 0.36^{+0.12}_{-0.10} Z_\odot$ (Rosati et al. 2004), and the surface brightness profile shows a discontinuity typical of cold fronts systems found in Chandra observations of low redshift clusters.

High angular resolution optical observations of the cluster with the Advanced Camera for Surveys (ACS; Ford et al. 1998) on the Hubble Space Telescope (HST) delivered unprecedented morphological information for the cluster galaxy populations as well as accurate photometry (Blakeslee et al. 2003b). Additional ground based photometry in the near-IR (Lidman et al. 2004) has been obtained with ISAAC[20] on the ESO Very Large Telescope (VLT), while Spitzer/IRAC observations of RDCS J1252.9-2927 (Stanford et al., in preparation) have made it possible to sample the rest-frame near-IR light of the cluster.

During the past few years we have carried out an extensive spectroscopic survey of RDCS J1252.9-2927 with the ESO VLT aimed at confirming a large fraction of the cluster population. This effort has yielded the most complete spectroscopic dataset so far on this cluster. In this paper we present the results from this ESO VLT/FORS[21] spectroscopic program in combination with ACS and VLT/ISAAC imaging data. This investigation is aimed at spectroscopically identifying the cluster galaxy populations in order to study the cluster dynamics from galaxy kinematics and the spectrophotometric properties of cluster members. Unless otherwise indicated, we assume a ΛCDM cosmology with $H_0 = 70$ km s^{-1} Mpc^{-1}, Ω_M=0.3, and Ω_Λ=0.7.

2. DATA SET AND DATA REDUCTION

2.1. *ACS and ground based imaging data*

Imaging observations of RDCS J1252.9-2927 in the optical have been carried out from the ground and from space. The ACS Wide Field Camera (WFC) on HST was used to obtain imaging with the F775W and F850LP filters (hereafter i_{775} and z_{850}, respectively), as part of a guaranteed time observation program (ID: 9290). Three orbits in i_{775} and five orbits in z_{850} were completed during May and June of 2002, distributed in a 2×2 mosaic pattern. The pointings overlapped about 1′, producing

[19] R_Δ is defined as the radius within which the mean density is Δ times the critical density of the universe at a given redshift.

[20] ISAAC stands for Infrared Spectrometer And Array Camera.

[21] FORS stands for FOcal Reducer and low dispersion Spectrograph.

exposures of 12 orbits in i_{775} and 20 orbits in z_{850} of the cluster core. A more detailed description of the ACS data on RDCS J1252.9-2927 is presented in Blakeslee et al. (2003b).

The ground based optical data were collected with FORS2 (Appenzeller & Rupprecht 1992) on the VLT, under ESO program 169.A-0458(A). The imaging data in the B-, V- and R-band[22] were obtained between March 1st and March 3rd, 2003, in very good seeing conditions. The co-added images have FWHM of $0\farcs66$ in the B-band, $0\farcs60$ in the V-band, and $0\farcs56$ in the R-band. The photometric calibration was done using the many Landolt field stars that were observed during the same nights. The images have been corrected for the instrumental response, bias and flat, in a standard manner. All images were photometrically aligned by using about one hundred bright but unsaturated stars in the field, which showed that the nights were indeed photometric. The images were astrometrically flattened and then finally co-added. As a further check, the same point sources used for the single images photo alignment were compared with FORS1 data acquired in May 1999, in the same B-, V-, and R-band. The mean differences in magnitude for the selected objects were -0.006±0.030 in B and 0.015±0.021 in V. In the R-band, a $V-R$ color term correction was found giving a mean difference of 0.002±0.032.

Ground based near-IR imaging observations of RDCS J1252.9-2927 were taken with ISAAC (Moorwood et al. 1999) on the ESO VLT and are described in Lidman et al. (2004). In the central regions of the cluster the data reach limiting magnitudes (5σ limit over a $0\farcs9$ diameter aperture) of 26.5 (AB) and 26.0 (AB) in the J- and K_s-bands, respectively. Since then, these data have been reprocessed to optimise image quality and to facilitate the computation of aperture corrections over the entire region covered by near-IR observations. The image quality in the reprocessed images varies from $0\farcs32$ to $0\farcs43$. The FWHM of point sources in a single tile is relatively constant and this makes it simpler to compute these aperture corrections, whereas the FWHM of point sources in the mosaic varies from tile to tile. Hence, rather than combining the images into a single mosaic, each tile was processed and analysed separately. This means that some objects appear multiple times as there is considerable overlap between different tiles. To handle this, we create catalogs for each tile and if a source is in more than one catalog, we average the result. We also compute an error from the different measurements and compare it to the error computed in SExtractor. If the source is in only one catalog, we use the error computed in SExtractor, otherwise we use either the SExtractor error or the error derived from the multiple measurements, whichever is the largest. Magnitudes are measured in 10 pixel diameter apertures ($\sim 1\farcs5$) and then corrected to larger apertures ($4''$diameter) using stars in each tile to compute aperture corrections.

2.2. *The photometric catalog*

A first multi-color catalog was built out of the above data sets. This catalog contained the photometry computed from the ISAAC mosaic as described in Lidman et al. (2004), without including the more recently reprocessed ISAAC photometry (see §2.1). We note that the original ISAAC photometry in Lidman et al. (2004) and the reprocessed one presented here (see below) are both in good agreement with each other, showing median differences of 0.002 in K_s and 0.007 in J_s with dispersion (σ) values of 0.064 and 0.033, respectively. The ACS, FORS2 and ISAAC images have different PSF FWHM. In order to avoid source blending due to a plain smoothing of the best seeing images to match the worst one, we adopt a different strategy to correct for these different PSFs. Using the same bright point-like sources as before (see §2.1), we construct growth curves. We then computed the aperture magnitude for all ground based and ACS images using an aperture of $1\farcs5$ diameter. An aperture correction from $1\farcs5$ to $4''$ diameter has been first computed using point sources, and then applied to all sources in the optical (ACS and FORS2) and near-IR (ISAAC) bands. Although this aperture correction is not strictly equivalent to the smoothing technique, we have checked, using the ACS data, that the two approaches give consistent results. The main advantage of using this approach instead of the smoothing one in this analysis is that we avoid significant source blending, especially in the central region of the cluster. This catalog was used to select candidates for spectroscopy while preparing most of the FORS2 masks.

The final photometric catalog published in this work and showing only spectroscopic cluster members is presented in Table 3. It contains the B-, V- and R-band photometry from FORS2, the i_{775}- and z_{850}-band photometry from ACS and the newly processed J_s and K_s ISAAC photometry, all of them as described above. The magnitudes are aperture corrected to a radius of $2''$ and are in the AB system (Oke 1974). To transform our photometry from the Vega system to the AB system, we recomputed AB corrections adopting the latest available Vega spectrum. The AB corrections for the B, V, R, J_s, J and K_s filters are -0.088, 0.052, 0.244, 0.968, 0.964 and 1.899 respectively. For the i_{775} and z_{850} filters we use AB correction values of 0.401 and 0.569, respectively (Blakeslee et al. 2003a). Values in Table 3 are corrected for galactic extinction. In order to take into account the galactic reddening in our analyses, we used the extinction maps of Schlegel et al. (1998). The corrections for the B-, V-, R-, i_{775}-, z_{850}-, J_s- and K_s- bands are 0.323, 0.248, 0.200, 0.145, 0.127, 0.067 and 0.027, respectively. Near-IR magnitudes presented in Fig. 2 and Fig. 14 are placed on the 2MASS J,K_s system (Carpenter 2001; Cohen et al. 2003). The corresponding transformation including the galactic reddening correction is given by:

$$(J-K_s)_{2MASS} = 1.038(J_s-K_s)_{ISAAC} - 0.048 \quad (1)$$

and

$$K_{s2MASS} = K_{sISAAC} - 0.028. \quad (2)$$

Unless otherwise stated, magnitudes presented throughout this paper are in the AB system.

2.3. *Photometric selection*

The above photometric information was used to compute photometric redshifts (Benítez 2000) for all the

[22] The B-, V- and R-band correspond to the B_BESSEL, V_BESSEL and R_SPECIAL (FORS2) or R_BESSEL (FORS 1) filters, respectively.

objects within the field of view (FoV) covered by the ISAAC mosaic (Lidman et al. 2004). This region was chosen in order to have optical ACS, optical FORS2 and near-IR ISAAC photometry available (see Fig. 1) to compute photometric redshifts. The set of galaxy templates for the photometric redshift calculations is presented in Benítez et al. (2004). Photometric redshifts together with color-color diagrams, were used to select candidate galaxies for spectroscopy. In our first spectroscopic campaign, galaxies with $K_s < 21$, $J_s - K_s < 2.1$ and $R - K_s > 3$ were targeted. This selection of the spectroscopic sample, indicated by the dashed lines in Fig. 2, minimizes the pollution by field galaxies, although it can be slightly biased against very red ($J_s - K_s \gtrsim 2.2$) cluster galaxies. This primary selection was complemented by another one on the $V - I$ vs $I - z_{850}$ plane in order to improve the selection of Balmer-Break and [OII] (λ3727) galaxies. The color tracks in Fig. 2 correspond to different evolutionary models, computed with the BC03 code (Bruzual & Charlot 2003), reproducing galaxy colors observed at $z = 1.237$ for different ages (from 1 to 5 Gyr). These models were employed together with the available multicolor information to guide the selection of targets. The orange lines represent single burst, solar metallicity models with burst durations of 0.5 Gyr (dot-dashed), 1.0 Gyr (dashed) and 2.0 Gyr (solid). The green lines represent exponentially declining star formation rate (SFR), solar metallicity models with characteristic times of $\tau = 1$ Gyr (dot-dashed), $\tau = 2$ Gyr (dashed) and $\tau = 5$ Gyr (solid). Model points, indicated by small open circles, are spaced by 0.5 Gyr in age and are in the ISAAC system. The blue line corresponds to 6 starburst galaxy templates from Kinney et al. (1996), hereafter KC templates, spanning the mean color excess range $0.0 < E(B-V) < 0.70$ and marked by open diamonds. In subsequent campaigns, we took advantage of photometric redshifts in the range $1.1 \lesssim z_{phot} \lesssim 1.3$ as they became available.

2.4. *VLT Spectroscopy*

A total of 15 masks were designed to be observed with FORS1 and FORS2. The data were obtained as part of an ESO Large Program (LP-166.A-0701) and two subsequent ESO proposals (69.A-0683 and 073.A-0832). The spectroscopic observations started in March 2000 and ended in April 2004. The first mask was observed in MOS mode with FORS1 and the 300I grism, whereas the other 14 masks were all observed with FORS2. Spectroscopy with FORS2 was carried out with both the 300I and 150I grisms. All the masks designed for MXU spectroscopy were observed with the 300I grism while all but one MOS mode observations were carried out with the 150I grism. The slit width for MXU observations was set to $1''$ while the MOS masks were designed with slit widths of $1\farcs4$ and $1\farcs7$. A summary of our spectroscopic observations is presented in Table 1 and more detailed information about the MOS and MXU modes as well as the grism characteristics are given in Demarco et al. (2005). From October 2002 on, the observations were carried out with a CCD mosaic composed of two 2k × 4k MIT/LL detectors with increased sensitivity ($\sim 30\%$) in the red, allowing us to obtain a better success rate compared to previous observations with the old FORS2 CCD. The data were reduced following the same procedure described in Demarco et al. (2005).

The efficient selection criteria of targets based on high quality imaging data has been a key piece in the success of our spectroscopic campaign. From a total number of 418 galaxies targeted with the 15 masks on FORS (see Table 1) we were able to obtain 282 (67%) redshifts, from which 227 (54%) correspond to secure measurements. We consider as cluster members all galaxies in the redshift range $1.22 < z < 1.25$, which corresponds to about $\pm 3\sigma_v$ around the cluster median velocity, where σ_v the cluster velocity dispersion (see §3.1). With this criterion 38 galaxies are classified as spectroscopic cluster members (see Table 2). The signal-to-noise ratio (SNR) per resolution element of cluster member galaxies in the range 3995 Å $< \lambda <$ 4085 Å rest-frame is observed to vary between ∼1 to ∼10. The redshifts and the corresponding error bars were obtained from the cross-correlation between the object and a template spectra as described in Demarco et al. (2005). In the cases of very poor SNR (SNR $\sim$ 1), the redshift was determined from the identification of [OII](λ3727) in the spectrum. In the case of passive galaxies, i.e., galaxies with an absorption line spectrum and no visible emission lines, the redshift was obtained from the cross-correlation technique as implemented in the RVSAO/XCSAO task in IRAF. In 23 objects (all with SNR > 2.5), CaII (H+K) features were identified as well. Yet, further spectroscopic classification, such as the one proposed by Dressler et al. (1999), was not possible with the current data due to the lack of an accurate measurement of the $H_\delta(\lambda 4102)$ feature. In general, we avoided observing the same object more than once, however, we obtained more than one redshift measurement for a small number of objects. This allowed us to estimate total (random+systematic) errors, yielding typical values of $\delta z \sim 12 \times 10^{-4}$. This value represents a more rigorous estimate of the redshift errors since the formal errors given by the cross-correlation (Tonry & Davis 1979) are known to be smaller than the true errors (e.g., Malumuth et al. 1992; Bardelli et al. 1994; Ellingson & Yee 1994; Quintana et al. 2000), and will be used in computations.

Those spectroscopic members for which FORS2 and ISAAC photometry is available (see table 3) are shown in Fig. 2. The red circles indicate spectroscopically confirmed non-emission line members, while blue triangles indicate emission line [OII](λ3727) confirmed members. The emission line object in the upper right side of the plot, marked with a square, corresponds to an X-ray source at the cluster redshift (ID=174; see Rosati et al. 2004; Martel et al. 2007). Note the unusual morphology of this source (see Fig. 5) with a faint red nucleus and an irregular, diffuse structure (also see Martel et al. 2007). In Table 4 we present the X-ray information corresponding to this source and other X-ray sources in the cluster FoV, for which a redshift measurement was obtained. The X-ray fluxes and rest frame X-ray luminosities in the soft ([0.5-2] keV) and hard ([2-10] keV) bands were extracted from the 190 ksec Chandra observations (Rosati et al. 2004) assuming a power law spectrum with photon index 1.4.

We estimated the success rate of our spectroscopic survey as the ratio of the number of objects with spectroscopic redshifts to the number of objects that were targeted for spectroscopy as a function of K_s (AB) mag-

nitude. The data were binned in $\Delta K_s = 0.5$ mag intervals and include only objects for which an estimate of their K_s magnitude is available, as shown in Fig. 3. Beside the overall success rate of our survey (solid histogram), the success rate for galaxies from the primary color-color selection (see section §2.1) is shown as the dashed histogram. A secondary color-color selection, defined by $I - z_{850} > 0.4$, $I - z_{850} < 0.85$, $V - I > 0.2$, $V - I < 1.2$, $V - I > [2.4\ (I - z_{850}) - 1.12]$ and $V - I < [7.0\ (I - z_{850}) - 2.3]$, was used to identify possible Balmer-Break and star-forming galaxies. The success rate of this additional cut in color-color space is shown in Fig. 3 as the dotted histogram. For galaxies within the primary color-color cut, the success rate decreases dramatically from about 85% to about 20% for galaxies with magnitudes fainter than 21.5 in K_s. At fainter magnitudes, the secondary color-color selection was more effective, as intended, at detecting sources with emission lines, increasing the number of galaxies per mask for which a redshift could easily be obtained. This secondary selection allowed us to obtain the redshift for 23 faint objects with $K_s > 21.5$. This corresponds to $\sim 12\%$ of all (no color restriction) targeted sources with $K_s > 21.5$. We note that all the sources within the $V - I$ vs $I - z_{850}$ cut with $K_s > 21.5$ and spectroscopic confirmation show some star-forming activity, while there are no targeted objects in this secondary color-color region with $K_s < 20.2$. The $V - I$ vs $I - z_{850}$ selection was intended to target star-forming galaxies at the cluster redshift, however only 8 confirmed members with [OII] were confirmed this way. In addition, no secure redshift was obtained for objects with $K_s < 20.5$ in this secondary color-color selecting area.

3. ANALYSIS

3.1. *Member selection and global properties*

The distribution in redshift space of the 226 galaxies with a secure estimate of their redshift is shown in Fig. 4. The bin size of the histogram is $\Delta z = 0.01$. We perform the adaptive-kernel method (hereafter DEDICA; Pisani 1993, 1996; Fadda et al. 1996; Girardi et al. 1996; Girardi & Mezzetti 2001) to search for the significant (>99% confidence level [c.l.]) peaks in the velocity distribution. This procedure detects RDCS J1252.9-2927 as the strongest peak at $z \sim 1.24$ (see Fig. 4) populated by 38 galaxies (hereafter referred to as cluster members), the largest number of spectroscopic members discovered so far in a cluster of galaxies at $z > 1$, distributed over a region of 1.8 Mpc in radius. Out of non-member galaxies, 170 and 20 are foreground and background galaxies, respectively (see Table 5). In particular, 33 foreground galaxies belong to a dense peak at $z \sim 0.74$. This structure, in the redshift range $0.70 < z < 0.79$ (see Table 5), has a 3σ-clipped mean redshift of $\bar{z} = 0.7429 \pm 0.0024$. The projected distribution of these sources extends over a region of about $5'.6$ a side in front of the cluster, showing a small concentration of about 10 galaxies at about $1'$ north from the cluster center. There is no X-ray extended emission associated to this structure indicating the lack of a hot gas component, although we note that there are 3 X-ray point sources at this redshift (see Table 4). This "group" may contribute to overestimate the cluster mass obtained by weak lensing analyses as discussed in Lombardi et al. (2005).

By applying the biweight estimator to the cluster members (Beers et al. 1990), we compute a mean cluster redshift of $\langle z \rangle = 1.2370 \pm 0.0004$. We estimate the line-of-sight (LOS) velocity dispersion, σ_v, by using the biweight estimator and applying the cosmological correction and the standard correction for velocity errors (Danese et al. 1980). We obtain $\sigma_v = 747^{+74}_{-84}$ km s^{-1}, where errors are estimated through a bootstrap technique. The spectroscopic information available on the cluster member sample is given in Table 2, while the ACS and ground based photometry of cluster members is presented in Table 3. Fig. 5 show $5'' \times 5''$ thumbnail images of the 37 spectroscopic cluster galaxies in the FoV of ACS together with the corresponding FORS spectrum. The left and middle panel correspond to the i_{775} (rest-frame U) and z_{850} (rest-frame B) data. The most prominent spectral features identified in each spectrum in the displayed wavelength range are indicated (see figure's caption for details). One cluster member, ID=7001, is out of the ACS FoV and its spectrum is shown in Fig. 6. For completeness, in Fig. 7 we show the optical spectrum of the 10 non cluster member AGN listed in Table 4. We do not present any analysis of the spectroscopic properties of these sources for being out of the scope of this work. Fig. 8 shows the distribution of RDCS J1252.9-2927 member galaxies in velocity and redshift space. The hatched area indicates the distribution of star-forming members. The secondary selection in $V - I$ vs $I - z_{850}$ allowed the discovery of 8 of the 17 cluster members with [OII](λ3727) emission lines. If we separate cluster members into two categories, passive and star-forming galaxies, there is no significant offset in median velocity between the two categories (see Fig. 8). The mean redshift of passive and emission line [OII] spectroscopic members is 1.2373 ± 0.0057 and 1.2369 ± 0.0054, respectively. The velocity dispersion of these populations are also consistent with each other and with the overall cluster velocity dispersion, within the uncertainties.

In Fig. 9 we show the rest-frame velocity vs. projected distance from the cluster center (lower panel). Hereafter we consider as cluster center the position of the X-ray center (Rosati et al. 2004): R.A.=$12^{\rm h}52^{\rm m}54.4^{\rm s}$, Dec.=$-29^{\rm o}27'17''.5$ (J2000.0). We also show the rest-frame velocities of the three brightest cluster members (IDs 291, 247, and 289, hereafter BCM1, BCM2, and BCM3). BCM1 is about 0.3 magnitudes brighter in K_s than BCM2 and BCM3 (upper panel). BCM1 and BCM3 are closely located near the position of the X-ray center and the peak of the galaxy distribution as recovered from 2D DEDICA method.

Assuming that RDCS J1252.9-2927 is in dynamical equilibrium (this will be discussed in §4) we can compute global virial quantities. Following the prescriptions of Girardi & Mezzetti (2001), we assume for the radius of the quasi-virialized region $R_{\rm vir} = R_{178} = 0.17 \times \sigma_v / H(z) = 1.61$ Mpc (see their equation 1 after introducing the scaling with $H(z)$; see also equation 8 of Carlberg et al. (1997) for R_{200}). Therefore, our spectroscopic catalog samples the whole cluster virialized region. We can compute the mass using the virial theorem (Limber & Mathews 1960; Girardi et al. 1998) under the assumption that the galaxies trace the total mass:

$$M = M_{\rm svir} - {\rm SPT}, \quad (3)$$

where

$$M_{\rm svir} = \left(\frac{3\pi}{2}\right) \times \left(\frac{\sigma_v^2\ {\rm R_{PV}}}{G}\right) \quad (4)$$

is the standard virial mass, $\rm R_{PV}$ a projected radius (equal to two times the harmonic radius), and SPT is the surface pressure term correction (The & White 1986). The estimate of σ_v is generally robust when computed within a large cluster region as shown in Fig. 10. At radii larger than 1.4 Mpc the profile becomes flatter, consistent with observations of low redshift clusters (Girardi et al. 1996; Fadda et al. 1996), suggesting that at large cluster radii any velocity anisotropy of cluster galaxies does not affect the value of $\sigma_v(< R)$. We thus consider the above global value, σ_v, as the cluster velocity dispersion. The value of $\rm R_{PV}$ depends on the size of the considered region, so that the computed mass increases (but not linearly) when increasing the considered region. Using the 37 galaxies within $\rm R_{vir}$ we obtain $\rm R_{PV} = 1.09 \pm 0.15$ Mpc. As for the SPT correction, we assume a correction factor of 20%, which we obtained by combining data on many clusters (e.g., Carlberg et al. 1997; Girardi et al. 1998). This leads to a virial mass $M(< {\rm R_{vir}} = 1.61\, Mpc) = 5.3^{+1.3}_{-1.4} \times 10^{14} M_\odot$.

We also used an alternative estimate of the virial mass (see eq. 13 of Girardi et al. 1998). This alternative estimate is based on the knowledge of the galaxy distribution and, in particular, a galaxy King-like distribution with parameters typical of nearby/medium-redshift clusters: a core radius $\rm R_{core} = 1/20 \times R_{vir}$ and a slope-parameter $\beta_{\rm fit} = 0.8$ (which gives a volume galaxy density at large radii as $r^{-3\beta_{\rm fit}} = r^{-2.4}$; Girardi & Mezzetti 2001). For the whole virialized region we obtain $\rm R_{PV} = 1.20$ Mpc where a 25% error is expected due to the fact that typical, rather than individual, galaxy distribution parameters are assumed. This leads to a virial mass $M(< {\rm R_{vir}} = 1.61$ Mpc$)=5.9^{+1.9}_{-2.0} \times 10^{14} M_\odot$, in good agreement with the above value.

To compare our result with the estimate obtained from the X-ray data for a smaller cluster region (Rosati et al. 2004) we assume that the cluster is described by a King-like mass distribution (see above) or, alternatively, a NFW profile where the mass-dependent concentration parameter is taken from Navarro et al. (1997) and rescaled by the factor $1 + z$ (Bullock et al. 2001; Dolag et al. 2004). We obtain $M_{\rm proj}(<$ R $= 0.536$ Mpc$)=(1.6-2.3)\times 10^{14} M_\odot$ in good agreement with that found by Rosati et al. (2004). This value is also in agreement with that found by Ettori et al. (2004) within a similar projected radius. Using the same mass distributions we also compute the projected mass within 1 Mpc for comparison with the weak lensing analysis by Lombardi et al. (2005). The cluster mass distribution is truncated at one virial radius or, alternatively, at two virial radii. The range of our results is $M_{\rm proj}(<$ R $= 1$ Mpc$)=(4.5-6.3)\times 10^{14} M_\odot$, in agreement with Lombardi et al. (2005) when taking into account the uncertainties on the mass values, although systematically lower.

By comparing the cluster X-ray luminosity (Rosati et al. 2004) and the cluster velocity dispersion derived above with the $L_X - \sigma_v$ relation presented in Fig. 2 of Xue & Wu (2000), we see a good agreement between the observed L_X and σ_v values in RDCS J1252.9-2927 and those from more than a hundred galaxy clusters selected from the literature. Using the best fit $\sigma_v - T$ relation from Xue & Wu (2000), without taking into account the fit errors, we obtain a temperature of $T = 3.89^{+0.61}_{-0.65}$ KeV for a velocity dispersion of $\sigma_v = 747^{+74}_{-84}$ km s^{-1}. This temperature is more than 3σ away from the observed value of $T_x = 6.0^{+0.7}_{-0.5}$ (Rosati et al. 2004), and still $\sim 0.5\sigma$ lower if the fit errors are considered.

3.2. *Substructures and projected distribution*

We analyze the velocity distribution to look for possible deviations from Gaussianity that could provide important signatures of complex dynamics. For the following tests the null hypothesis is that the velocity distribution is a single Gaussian.

We compute three shape estimators, i.e. the kurtosis, the skewness, and the scaled tail index (see, e.g., Beers et al. 1991). The value of the normalized kurtosis (-0.754) shows evidence that the velocity distribution differs from a Gaussian, being lighter-tailed, with a c.l. of $\sim 95-99\%$ (see Table 2 of Bird & Beers 1993).

Then we investigate the presence of gaps in the distribution. A weighted gap in the space of the ordered velocities is defined as the difference between two contiguous velocities, weighted by the location of these velocities with respect to the middle of the data. We obtain values for these gaps relative to their average size, precisely the midmean of the weighted-gap distribution. We look for normalized gaps larger than 2.25 since in random draws of a Gaussian distribution they arise at most in about 3% of the cases, independent of the sample size (Wainer & Schacht 1978; Beers et al. 1991). One significant gap in the ordered velocity dataset is detected dividing the dataset in two subsets containing 26 and 12 galaxies from "low" to "high" velocities (hereafter WGAP1 and WGAP2 groups; see top panel in Fig. 9). We compare these two subsets applying the the 2D Kolmogorov-Smirnov tests to the galaxy positions (2DKS-test; see Fasano & Franceschini 1987), as implemented by Press et al. (1992). In spite of the modest statistics we find a marginal significance (89% c.l.) for the difference. The "high" velocity WGAP2 group is located towards East with respect to the "low" velocity group WGAP1 (see Fig. 11).

To further investigate the substructure membership we use the Kaye's mixture model (KMM) test as implemented by Ashman et al. (1994). The KMM algorithm fits a user-specified number of Gaussian distributions to a dataset and assesses the improvement of that fit over a single Gaussian. In addition, it provides the maximum-likelihood estimate of the unknown n-mode Gaussians and an assignment of objects into groups. However, one of the major uncertainties of this method is the optimal choice of the number of groups for the partition. Using the results of the gap analysis we decide to fit two groups and determine the first guess for the group partition. We do not find any group partition which provides a significantly better description of the velocity distribution with respect to a single Gaussian in the 1D analysis. However, since the 3D diagnostic is, in general, the most sensitive indicator of the presence of

substructure (e.g., Pinkney et al. 1996)), we apply the 3D version of the KMM software using simultaneously galaxy velocity and positions. In the 3D case we find that a two-group partition of 26 and 12 galaxies (hereafter KMM1 and KMM2 groups) better describes the velocity distribution at the 95.6% c.l., according to the likelihood ratio test (see Fig. 9). These groups are located at -354 ± 159 km s^{-1} and 845 ± 182 km s^{-1} with respect to the cluster mean velocity, and to the West and East of the cluster center, respectively. The velocity dispersions of KMM1 and KMM2 are estimated to be 486^{+47}_{-85} km s^{-1} and 426^{+57}_{-105} km s^{-1}, respectively. This partition corresponds to that indicated by the weighted gap analysis with the difference of two galaxies (see Fig. 11). These velocity dispersion values are less robust than the overall cluster velocity dispersion due to uncertainties in the membership of galaxies to KMM1 and KMM2, although the detection and location of the substructures are reliable. By assuming dynamical equilibrium and theoretically estimating R_{PV} for each KMM group as done in section §3.1 for the whole cluster, we obtain masses of $M(< \mathrm{R}_{200} = 1.05$ Mpc$)=1.6^{+0.5}_{-0.7} \times 10^{14} M_{\odot}$ for KMM1 and $M(< \mathrm{R}_{200} = 0.92$ Mpc$)=1.1^{+0.4}_{-0.6} \times 10^{14} M_{\odot}$ for KMM2. We note that the sum of these masses is about the half of the cluster mass; we will come back to this in section §4.

To look for further evidence that RDCS J1252.9-2927 is still not completely dynamically relaxed, we analyzed the velocity of the brightest galaxy BCM1. In fact, since its location coincides with the X–ray center, one expects that BCM1 is at the center of the cluster potential and thus at the center of the velocity distribution. We find that BCM1 shows evidence of peculiarity at the $> 95\%$ c.l. according to the Indicator test by Gebhardt & Beers (1991), while it lies very close to the peak of the velocity distributions of KMM1 (see Fig. 9). We finally note that the close couple of luminous galaxies BCM1 and BCM3 are separated by ~ 550 km s^{-1} in rest-frame velocity and, as first noted by Blakeslee et al. (2003b), in interaction as supported by Rettura et al. (2006) who show signs of interaction in the form of an S-shaped residual after galaxy subtraction linking the two galaxy centers. In addition, BCM2 is near the edge of the coma-like structure of the X-ray surface brightness reported by Rosati et al. (2004).

In Fig. 1 we show the projected distribution of RDCS J1252.9-2927 spectroscopic members on the plane of the sky[23]. Circles are passive galaxies and triangles are emission line members. The overall shape of the distribution of confirmed galaxy members is clearly elongated in the East-West direction, as previously reported in Toft et al. (2004) based on the K_s-band light distribution of photometric cluster members, although more uniform than that of other high redshift clusters (see, e.g., Gioia et al. 1999; Demarco et al. 2005; Girardi et al. 2005). The projected distribution of the KMM groups is consistent with the overall cluster galaxy distribution, suggesting that the observed elongation is caused by the merger of both groups.

3.3. *Spectral properties of cluster galaxies*

[23] The background image corresponds to the ACS combined i_{775}- and z_{850}-bands data.

The spectral features of cluster members indicate the stellar content of galaxies in RDCS J1252.9-2927 and for some members, the [OII] line was used to estimate their SFR. The spectrum of each one of the 37 cluster members in the ACS FoV is shown in the right panel of Fig. 5 while the spectrum of the spectroscopic member outside the FoV of ACS is shown in Fig. 6. The flux calibration of the spectra is not accurate in the very red. Following Homeier et al. (2005), we estimated the SFR for the 17 member galaxies with [OII]λ3727 in their spectra. We measured the integrated flux of the [OII] line using the bandpass defined in Tran et al. (2003). In two cases with a very low signal-to-noise continuum (IDs 174 and 9000), the [OII] line flux was underestimated due to sky over-subtraction. Previous to these measurements, all the spectra were doppler corrected by using the task DOPCOR in IRAF. These line fluxes were converted into $L_{[OII]}$ luminosities by assuming that all these galaxies are at the same distance, corresponding to the median redshift of the cluster ($z = 1.237$). SFRs were derived from the $L_{[OII]}$ values by following the prescription of Kewley et al. (2004):

$$SFR_{[OII]}(M_{\odot}\, yr^{-1}) = (6.58 \pm 1.65) \times 10^{-42} L_{[OII]}(ergs\, s^{-1}), \quad (5)$$

which takes into account the mean reddening corrected [OII]/$H\alpha$ $= (1.2 \pm 0.3)$ ratio from the Nearby Field Galaxies Survey (Jansen et al. 2000). We do not correct for metallicity and dust extinction effects on the [OII] flux measurements, therefore our SFRs estimates should be considered as lower limits to the true SFR. The [OII] line can be considerably affected by dust absorption, giving SFRs one or two orders of magnitude lower than SFRs unaffected by dust at 15 μm (Duc et al. 2002; Coia et al. 2005). A more robust SFR indicator is the $H\alpha$ line (see, e.g., Charlot & Longhetti 2001), however, at the cluster redshift this line is not observed in the optical. IR luminosities can also be used as reasonable tracers of SFR in galaxies (Kewley et al. 2002; Calzetti et al. 2005). Integrated [OII] line flux measurements, [OII] equivalent widths and SFR values derived from $L_{[OII]}$ of star-forming cluster galaxies are given in Table 2. Error bars are estimated by taking into account the RMS fluctuation in flux within the two sidebands at both sides of the [OII] bandpass, as defined in Tran et al. (2003). The median value of the derived SFRs is 0.7 $M_{\odot}$ yr^{-1}, with a few galaxies reaching SFRs greater than 2.0 $M_{\odot}$ yr^{-1} (see Table 2). Typical errors in the SFR are about 47%, reaching values greater than 60% in a couple of cases. Our median SFR value is consistent within the errors with the mean SFR ($0.17 \pm 0.02\ h_{100}^{-2} M_{\odot}\, yr^{-1}$) derived by Balogh et al. (1998) from the [OII](λ3727) emission line for cluster galaxies at large (R_{200} $\sim$1.5-2 h_{100}^{-1} Mpc) clustercentric radii and in the redshift range $0.18 < z < 0.55$. These values are, however, much lower than that of $\sim 3 h_{100}^{-2} M_{\odot}\, yr^{-1}$ derived for cluster galaxies at $z \sim 0.75$ from $H\alpha$ (Finn et al. 2005), and even lower than the $H\alpha$ SFR of field galaxies at $z \sim 1$ (Glazebrook et al. 1999; Doherty et al. 2004). In spite of the little or no evolution in SFR between the sample of Balogh et al. (1998) and our sample, this comparison can be affected by dust (see section §3.5) and metallicity effects on the [OII] line, therefore, a fair comparison requires Hα and

IR measurements of the true SFR in order to quantify the amount of evolution. The lack of SFR measurements in cluster galaxies at $z > 1$ and the limitations of our data prevent us from studying the evolution of the SFR in clusters from $z \sim 1.3$ down to the local universe.

The co-added spectrum of the 17 star-forming cluster members is shown in Fig. 12. The presence of young stellar populations (A- and F-type stars) is inferred by the detection of Balmer absorption lines. Most of these galaxies present irregular, extended features, most likely the home for the young stars and where most of the star-forming activity takes place. In particular, we note that object ID=6306, with an $i_{775}-z_{850} = 0.85$ color, presents a spectrum showing a young post-starburst component (prominent Balmer absorption features) in addition to an on-going star-formation activity ([OII] in emission; see Fig. 5). This galaxy can therefore be classified as an E+A+[OII] galaxy (e.g., Demarco et al. 2005). We also note that ID=7001 has similar characteristics.

On the other hand, there is an indication of the remaining traces of the latest episode of star formation in some of the massive, passive early-type galaxies in the cluster. Due to the lower signal-to-noise ratio of our spectra at wavelengths larger than about 9000 Å, an accurate measurement of the EW of the H_δ absorption feature is not possible. Therefore, an identification of post-starburst stellar populations in individual early-type cluster galaxies cannot be properly done with the existing data. However, by co-adding the spectra of the 10 brightest passive galaxies in RDCS J1252.9-2927 (Rosati 2004), a prominent H_δ feature emerges and some other higher-order Balmer features become visible. In Fig. 13 we show the co-added spectra of the 10 (top panel) and 20 (bottom panel) brightest (in K_s) passive cluster members. An increase of the H_δ (about 36%) and other Balmer absorption features equivalent width can be seen when including fainter early-type members. We note that the H_δ absorption line can indeed clearly be seen in the individual spectrum of the first, second, third and fifth brightest cluster members (in z_{850}) after a total integration time of 24 hr (Holden et al. 2005). A more detailed quantitative analysis of this observation is underway (Rosati et al., in preparation).

3.4. *Color-magnitude distribution of cluster galaxies*

Lidman et al. (2004) included all galaxies within 20″ (0.17 Mpc) of the cluster center in their $J-K_s$ versus K_s color-magnitude diagram of RDCS J1252.9-2927. In this paper, we produce a second color-magnitude diagram using the reprocessed ISAAC data, but here we only use galaxies with measured redshifts. The stellar symbols in Fig. 14 represent spectroscopically confirmed stars. Crosses are non-cluster members, i.e., objects with redshift $z \leq 1.22$ or $z \geq 1.25$. Filled circles are cluster members without detectable [OII](λ3727) emission. The filled triangles, on the other hand, are cluster members with [OII] emission. The dotted red line is the fit published in Lidman et al. (2004). At that time, only a limited number of redshifts were available, so, to limit the effect of contamination from non-cluster members, only galaxies within 20″ of the cluster center and within the blue rectangle in Fig. 14 were used in fitting the CM relation. Here, only cluster members without [OII] are used in the fit, which is shown as the solid red line. No other restrictions are used. The slope and scatter about the two fits are listed in Table 6 and are calculated as explained in Lidman et al. (2004). Within the statistical error both fits are the same. Therefore, as in Lidman et al. (2004), the values here obtained for the slope and scatter about the new fit imply that passive galaxies in RDCS J1252.9-2927 are home to old stellar populations with a mean age of ~ 3 Gyr (according to solar-metallicity, single stellar population models of Bruzual & Charlot 2003). This corresponds to a formation redshift of $z_* \sim 3$ for the bulk of the stars in these galaxies. The values of the slope and scatter presented in this work were obtained using all passive members up to more than 0.5 Mpc from the center, without restricting ourselves to the central 40″ (diameter; ~ 0.3 Mpc) as in Lidman et al. (2004), and are consistent with measurements in lower redshift clusters (Stanford et al. 1998).

Looking at morphology and colors, we do not see any indication of the existence of a population of S0 galaxies with a bluer (with respect to elliptical galaxies) color-magnitude sequence, in contrast to what has been observed in the galaxy cluster RDCS J0910+5422 at $z = 1.1$ (Mei et al. 2006a). As pointed out by Mei et al. (2006a), the existence of a bluer color-magnitude sequence of the S0 galaxies with respect to the elliptical galaxies in RDCS J0910+5422 could be the result of a still forming cluster, with these bluer S0 galaxies being part of a group infalling from the field onto a more evolved red cluster population. RDCS J1252.9-2927 presents a more evolved structure having a more evolved early-type population, with elliptical and S0 galaxies distributed over a common red sequence.

From Fig. 14, we observe that all spectroscopic members without detectable [OII](λ3727) emission have magnitudes brighter than $K_s = 21.5$, corresponding to $\sim K_s^* + 1.5$ ($K^*_{s,rest} = -24$) at z=1.237 (Strazzullo et al. 2006). This indicates the spectroscopic limit for galaxies with no detectable [OII] emission, while star-forming objects can be confirmed down to $K_s \sim 24$ ($\sim K_s^* + 4$). Star-formation is active in faint ($> K_s^* + 1.5$), less massive objects with only a few star-forming members with $K_s < 21.5$. One of these objects is ID=174, the cluster member AGN, which is also the reddest ($J - K_s \sim 1.3$) member in Fig. 14. The other star-forming galaxies (IDs 309, 339, 445 and 345) are well located within the near-IR CM relation.

If the fit of the cluster red sequence were restricted to cluster members brighter than $\sim K_s^*$, the resulting slope would be considerably flatter. This steepening of the slope when including fainter passive members could be due in part to the selection effect imposed by our inability of detecting passive members below the spectroscopic limit of our survey. Star-forming objects, having bluer colors, would contribute to produce a fit slope steeper than the one obtained if only more luminous and passive (redder) objects were considered. The selection effect caused by the spectroscopic limit of passive galaxies prevent us from properly investigating the faint end of the red cluster sequence, in particular its build-up at this high redshift. As indicated by Tanaka et al. (2005), the faint end of the red cluster sequence seems to be still under construction at $z \sim 0.8$. The observed deficiency of galaxies at the faint end in $z \sim 0.8$ clusters suggests that

part of the blue cluster galaxy population may well be the progenitors of the present day faint end population of the red sequence (De Lucia et al. 2004), as supported by recent high-resolution simulations (De Lucia et al. 2006) which, however, find it difficult to reproduce such tight color sequences at this redshift. More observations are needed to firmly establish this.

In Fig. 15 we show the color-magnitude diagram using the ACS filters. Here we include spectroscopically confirmed passive (red circles) and emission line (blue triangles) cluster members. Squares are objects with photometric redshifts (Benítez 2000) in the range $1.12 < z_{phot} < 1.35$ and within an aperture of 1 Mpc centered in the cluster (we excluded known spectroscopic non-members). These photometric redshifts have been corrected for a systematic deviation $\langle z_{spec} - z_{phot} \rangle = 0.17$. The red sequence can be traced down to a magnitude of $z_{850} = 24.5$ in these ACS bandpasses, vanishing quickly at fainter magnitudes. However, this cutoff magnitude can be affected by uncertainties in the photometric redshifts.

3.5. *Color-Color distribution of cluster galaxies*

The distribution of spectroscopic cluster members in color-color space is shown in Fig. 16. Red filled circles correspond to passive galaxies and blue filled triangles to [OII] emission galaxies. The open square indicates the only known AGN at the cluster redshift (ID=174; see section §2.4). A clear separation between passive and star-forming galaxies is seen, with the exception of three emission line galaxies, IDs 309, 345 and 445, which are located in the locus of passive members. These galaxies are also located in the red sequence of both $i_{775} - z_{850}$ vs z_{850} and $J_s - K_s$ vs K_s color-magnitude diagrams. The estimated SFR for these objects varies from 0.5 to 2.9 $M_\odot$ yr^{-1} (see table 2), although the error bars are large.

The shape of the continuum and the [OII](λ3727) line of IDs 345 and 445 are consistent with their red ($i_{775} - z_{850} \sim 0.95$, $J_s - K_s \sim 0.8$) colors and star-forming nature. ID=309, on the other hand, shows a rather flat continuum up to 9000 Å suggesting that we are looking at a dust-rich star-forming galaxy. The increase in number of sky lines, a relatively low SNR (SNR$\sim$3) and therefore a poorer extraction of the continuum beyond 8500 Å may have caused the flattening of the spectrum towards red wavelengths, in contrast to its red $i_{775} - z_{850}$ color. These galaxies are located within an annulus spanning from ~ 0.4 to ~ 0.8 Mpc in radius from the two central bright elliptical galaxies, which corresponds to $R/R_{200} \simeq 0.43 - 0.85$ in Postman et al. (2005). The morphology-radius relation obtained for a composite cluster sample of 7 galaxy clusters (including RDCS J1252.9-2927) at $z \sim 1$ observed with the ACS (Postman et al. 2005) shows that these objects are in a region where the average fraction of early-type galaxies ($\sim$39%) is about half of that of late-type galaxies ($\sim$61%). Red, possibly obscured, star-forming galaxies in clusters at low redshift have already been reported in the literature (Duc et al. 2002; Coia et al. 2005; Wolf et al. 2005; Popesso et al. 2006), while similar objects have also been observed at redshift $z > 0.8$ (van Dokkum & Stanford 2003; Demarco et al. 2005) and in the field up to $z \sim 2$ (Franzetti et al. 2006).

Finally, we note that the two other [OII] galaxies with colors $J_s - K_s > 0.9$ (IDs 248 and 339), excluding the X-ray source ID=174, can be starburst galaxies with $E(B-V) > 0.39$ (Kinney et al. 1996) as shown in Fig. 2 and Fig. 16, while the colors of the rest of the member star-forming galaxies can be reproduced by galaxies with $Z = Z_\odot$ and exponential SFRs (see Fig. 2 and Fig. 16). We also note that ID=339, with an elongated and irregular morphology, has a blue color in the optical ($i_{775} - z_{850} = 0.64$) but a red color in the near-IR ($J_s - K_s = 0.92$), suggesting that this star-forming galaxy has a significant amount of dust obscuring an important fraction of its "blue" light.

3.6. *Morphologies and spectrophotometric properties*

The morphological T-type class (de Vaucouleurs et al. 1976) as determined by Postman et al. (2005) for cluster members is given in Table 3. T-type values range from -5 for elliptical galaxies to 8 for disk/irregular galaxies. A value of -2 corresponds to S0 galaxies, values between -1 and 1 are assigned to morphologies between S0 and Sa and a value of 6 corresponds to an Sd morphology. In general, we observe the well-known correlation between morphology and SFR. Most of the star-forming objects in this cluster have irregular morphologies or irregular disky structures where star formation is taking place. One of these, ID=6301, can be a possible merger with two compact bright regions surrounded by a gas envelope. Another one, ID=619, shows clear spiral-arm features and could also be a merger. In contrast, most of the passive galaxies have morphologies typical of elliptical and S0 galaxies, with no on-going star formation due to their poor or zero gas content.

Passive early-type galaxies have red colors while star-forming late-type galaxies are blue, (see Fig. 14), a manifestation of the well-known correlation between stellar populations and morphology. However, in Fig. 14 we observe that a few star-forming galaxies have colors as red as the passive cluster members in the red sequence. One of these sources is ID=174, the confirmed cluster member X-ray source (see section §2.4), with a $J - K_s$ color about 0.35 magnitudes redder than the average color of red sequence members and whose star-formation can be due to the central, dust-obscured AGN. The other red, star-forming sources in the red sequence of Fig. 14 are IDs 248, 309, 339, 345 and 445. While IDs 248 and 339 show an irregular disky appearance, IDs 309, 345 and 445 are characterized by a prominent central bulge surrounded by a diffuse disk structure (see Fig. 5). We observe that IDs 309, 345 and 445 also have red $i_{775} - z_{850}$ colors ($i_{775} - z_{850} \sim 0.95$) that locate them in the locus occupied by red, passive galaxies in Fig. 16. While the star-formation in these sources must be occurring in their gas-rich regions, their red colors should arise from the combined effect of old stellar populations and dust reddening. One interesting object is ID=726. This source is very compact and also shows a clear [OII] emission line in its spectrum, as shown in Fig. 5. In addition, its blue color (see table 3) is consistent with the flat continuum of its spectrum. With a magnitude $M_B = -20.74$, this object can be considered in the class of luminous compact blue galaxies reported in the literature (e.g., Werk et al. 2004; Hammer et al. 2005; Noeske et al. 2006). This object has a very compact morphology (it is similar to the PSF), hence Postman et al. did not provide a

morphological T-type for it (see Table 3).

4. DISCUSSION AND CONCLUSIONS

We have combined an extensive multi-wavelength dataset to provide a comprehensive picture of the X-ray luminous, massive cluster of galaxies RDCS J1252.9-2927 at $z = 1.237$. ACS data provides detailed morphological information of galaxies and accurate photometry. By using FORS on the VLT, we were able to spectroscopically confirm 38 cluster members. The overall projected distribution of spectroscopic cluster members presents a clear elongation in the East-West direction, consistent with the elongation observed in the gas (Rosati et al. 2004) and dark matter (Lombardi et al. 2005) distributions. As pointed out by Rosati et al. (2004), the X-ray surface brightness distribution of the core of RDCS J1252.9-2927 has a comet-like shape, resembling a cold front in low redshift clusters (e.g., Mazzotta et al. 2001) produced by a merging process of a subclump in the East-West direction.

The velocity dispersion σ_v of the cluster members is 747^{+74}_{-84} km s^{-1}. Assuming virial equilibrium we find that the virial mass is $\sim 5 \times 10^{14} M_\odot$. Our mass estimates are also consistent with the mass estimates from the X-ray data (Rosati et al. 2004; Ettori et al. 2004) and the weak-lensing mass map (Lombardi et al. 2005). Our value of σ_v is also consistent with the $L_x - \sigma_v$ relation for clusters, but turns out to be in disagreement with the $\sigma_v - T_x$ relation (Xue & Wu 2000). While the observed L_x and T_x values follow the $L_x - T_x$ relation of Xue & Wu (2000), the deviation from the $\sigma_v - T_x$ relation may be the signature of a galaxy distribution still not in a fully dynamical equilibrium.

Indeed, RDCS J1252.9-2927 has not yet attained an equilibrium state as suggested by a more detailed analysis of the velocity distribution of cluster members. RDCS J1252.9-2927 is the most distant cluster for which a substructure analysis can be attempted. The kurtosis (Beers et al. 1991) indicates that the velocity distribution of cluster members is not Gaussian at the $> 95\%$ c.l.. The weighted gap procedure finds a significant gap in the space of ordered velocities and a 3D KMM study detects substructure in velocity at the 95.6% c.l.. The substructure is composed by two groups of galaxies which are mostly distributed in a East-West direction on the sky, as shown in Fig. 11. Twenty-six galaxies are confirmed in the "low-velocity" group (centered at -354 ± 159 km s^{-1} with respect to the cluster mean velocity) with a velocity dispersion of 486^{+47}_{-85} km s^{-1} and located mostly to the West of the cluster center. Twelve galaxies are confirmed in the "high-velocity" group (centered at 845 ± 182 km s^{-1} with respect to the cluster mean velocity) with a velocity dispersion of 426^{+57}_{-105} km s^{-1} and to the East of the cluster center.

The existence of the above correlations between positions and velocities of cluster galaxies is a footprint of real substructure. In particular, it suggests that RDCS J1252.9-2927 is forming via a merger along the E-W direction. Going further, we estimate the mass contained within R_{200} for each group (see section §3.2). The sum of these two masses is about half of the cluster mass, and it should be considered as a lower limit of the latter. These groups may be the virialized cores of two larger systems, with possibly $R \gtrsim 2\ R_{200}$ and larger velocity dispersions. The fact that the X-ray emission presents a single peak, although elongated in morphology, indicates that these two groups have already started virialization. This is supported by the degree of merging of the group members as shown in Fig. 11. Our substructure analysis has thus found the remnant traces of two sub-clusters merging parallel to the sky as suggested by the lack of a clear bimodal distribution of the overall cluster velocity field.

By using the $\sigma_v - T_x$ relation for groups in Xue & Wu (2000), we find that both groups have velocity dispersions consistent with the same temperature of $\sim$2 keV. This temperature is inconsistent with the overall cluster temperature of $\sim$6 keV, yet it corresponds to about half the total cluster X-ray luminosity according to the group $L_x - T_x$ relation (Xue & Wu 2000). Thus, both groups seem to be independently virialized and the sum of their X-ray luminosities would be approximately equal to the total cluster L_x. The temperature value of the overall ICM would have been the consequence of gas particle interactions at an earlier epoch during the merger of the groups. By the epoch of observation, the overall ICM is consistent with being in an almost isothermal hydrostatic equilibrium in a single potential well, while the groups are fossilized in the velocity dispersion substructure. The agreement of the observed cluster L_x and T_x values with the $L_x - T_x$ relation and the deviation of the cluster values from the $\sigma_v - T_x$ relation show that the gas in the cluster is more relaxed than the galaxies. Since galaxies behave more like dark matter particles, they still need more time to reach the $\sigma_v - T_x$ relation expected for the cluster. In summary, the elongated shape of the cluster structure (in all its components: baryons and dark matter) and the substructure in velocity leads us to conclude that RDCS J1252.9-2927 has not yet reached a final virial state: we may be seeing gravitational collapse of two or more sub-clusters along a high density filament.

In this work we present an improved analysis of the near-IR color-magnitude diagram of galaxies in RDCS J1252.9-2927, based on ISAAC data which have been reprocessed in order to optimise image quality and obtain more accurate photometry over the entire region covered by near-IR observations. We use only spectroscopic cluster members and the analysis extends over more than 1 Mpc from the cluster center, in contrast to previous studies. The fit to the red cluster sequence, including all passive spectroscopic members, yields slope and scatter values consistent with the result of Blakeslee et al. (2003b) and Lidman et al. (2004). This indicates that the early-type galaxy cluster population has formed most of its stellar content at $z \sim 3$ and has passively evolved since down to the epoch of observation, as previously shown for $z \sim 1$ clusters (van Dokkum & Stanford 2003; Blakeslee et al. 2003b; Lidman et al. 2004; Blakeslee et al. 2006; Mei et al. 2006a,b). If we restrict the fit of the cluster red sequence to spectroscopic members brighter than $\sim K_s^*$, the slope becomes significantly shallower, suggesting that the CM relation of passively evolving systems seems to be non-linear, although this could also be due to a selection effect. The [OII] emission line observed in 17 cluster galaxies allows us to estimate a lower limit value of $0.7 M_\odot$ yr^{-1} for the median SFR in the cluster. None of the spectroscopic members brighter than $K_s^* + 1.5$ but

a few show ongoing star formation. Three star-forming galaxies have colors as red as galaxies in the red cluster sequence of both $i_{775}-z_{850}$ vs. z_{850} (optical) and J_s-K_s vs. K_s (near-IR) color-magnitude diagrams. One additional star-forming galaxy is observed in the red sequence of the near-IR color-magnitude diagram but not in the optical one, suggesting that we are observing a dust-rich galaxy. Finally, a highly obscured AGN is observed to be a spectroscopic member of RDCS J1252.9-2927, and we do not find any evidence of "red mergers" like those observed in the galaxy cluster MS1054 (van Dokkum et al. 1999).

The projected distribution of passive and star-forming galaxies in this cluster shows a spectrum-density relation that qualitatively resembles the observed morphology-density relation at $z \sim 1$ (Smith et al. 2005; Postman et al. 2005): passive, early-type galaxies dominate the cluster core while star-forming, late-type galaxies are found in the outskirts of the cluster. From a dynamical point of view, we are witnessing hierarchical structure formation: old and massive galaxies are already in place, as does the morphology/spectrum-density relation, but, at the same time, we are observing a merger of two large groups of galaxies into a more massive structure.

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TABLE 1
A summary of the VLT spectroscopic data obtained on RDCS J1252.9-2927. Grism/Filter corresponds to the grism/order sorting filter combination used and the exposure times are given in units of hours. The (*) indicates observations carried out with the mosaic of two 2k × 4k MIT CCDs.

Mask name	Date	Telescope	Instrument	Grism/Filter	Exp. time	Mask
m1	Mar. 2001	UT2	FORS2	300I/OG590	4.9	MXU
m2	Mar. 2001	UT2	FORS2	300I/OG590	4.0	MXU
m3	Mar. 2001	UT2	FORS2	300I/OG590	0.9	MXU
m3	Apr. 2001	UT2	FORS2	300I/none	4.5	MXU
m4	Mar. 2001	UT2	FORS2	300I/OG590	1.7	MOS
m4	Mar. 2001	UT2	FORS2	150I/GG435	1.4	MOS
m5	Mar. 2001	UT2	FORS2	150I/GG435	1.3	MOS
m6	Apr. 2001	UT2	FORS2	300I/none	4.5	MXU
m7	Apr. 2001	UT2	FORS2	150I/none	2.0	MOS
m8	Apr. 2001	UT2	FORS2	150I/none	2.0	MOS
m9	Apr. 2001	UT2	FORS2	150I/none	1.0	MOS
m10	Apr. 2001	UT2	FORS2	300I/none	2.5	MXU
m11	Oct. 2002	UT4	FORS2*	300I/none	6.5	MXU
m12	Feb. 2003	UT4	FORS2*	300I/none	3.7	MXU
m13	Feb. 2003	UT4	FORS2*	300I/none	3.0	MXU
m14	Apr. 2004	UT4	FORS2*	300I/none	4.7	MXU
m15	Apr. 2004	UT4	FORS2*	300I/none	4.7	MXU

TABLE 2

SPECTROSCOPICALLY-CONFIRMED CLUSTER MEMBERS. R.A. AND DEC. (J2000) COORDINATES ARE IN (H:M:S) AND (^{o}:$'$:$''$), RESPECTIVELY. COLUMN 4 SHOWS THE REDSHIFT AND ITS FORMAL ERROR AS OBTAINED FROM THE CROSS-CORRELATION. COLUMN 5 SHOWS THE CORRELATION COEFFICIENT, R, AS DEFINED IN TONRY & DAVIS (1979). COLUMN 6 INDICATES THE MOST PROMINENT SPECTRAL FEATURES IDENTIFIED IN THE SPECTRUM[a]. COLUMN 7 SHOWS THE EMISSION LINE FLAG ASSIGNED TO THE OBJECT. A VALUE OF 0 CORRESPONDS TO GALAXIES WITHOUT EMISSION LINES AND A VALUE OF 1 CORRESPONDS TO GALAXIES WITH NARROW EMISSION LINES. THE LAST THREE COLUMNS ARE THE INTEGRATED [OII] LINE FLUX (IN $\times 10^{-17}$ $erg\ s^{-1}\ \ cm{-}2$), THE [OII] LINE EQUIVALENT WIDTH (EW; IN Å) AND THE STAR-FORMATION RATE (SFR; IN $M_\odot\ yr^{-1}$) RESPECTIVELY.

ID	R.A.	DEC.	z	R	Spec. Feat.[a]	E. L.	Flux [OII]	EW [OII]	SFR
137	12:52:57.65	-29:28:07.5	1.2475±0.0004	2.37	MgI;CaII(H+K);D4000;H_δ	0	...	...	...
149	12:52:49.65	-29:28:03.7	1.2382±0.0003	3.45	MgI;CaII(H+K);D4000;H_δ	0	...	...	...
174	12:52:49.82	-29:27:54.9	1.2382±0.0002	4.73	[OII]	1	1.25±0.04	-596±528	0.7±0.2
205	12:52:51.98	-29:27:46.2	1.2318±0.0003	4.11	MgI;CaII(H+K);D4000;CaI	0	...	...	...
206	12:52:48.74	-29:27:45.3	1.2353±0.0003	3.11	H8;MgI;H6;CaII(H+K);D4000	0	...	...	...
247	12:52:51.15	-29:27:31.4	1.2351±0.0002	5.79	H8;H6;CaII(H+K);D4000;H_δ;CaI	0	...	...	...
248	12:52:57.39	-29:27:32.1	1.2322±0.0001	4.91	[OII];CaII(H);CaI	1	0.54±0.20	-63±24	0.3±0.2
265	12:52:57.64	-29:27:29.8	1.2358±0.0004	1.98	H8;CaII(K);D4000	0	...	...	...
282	12:52:54.42	-29:27:23.7	1.2472±0.0004	1.99	CaII(H+K);D4000;CaI	0	...	...	...
289	12:52:54.54	-29:27:17.1	1.2378±0.0006	2.23	H9;D4000;CaI	0	...	...	...
291	12:52:54.41	-29:27:17.7	1.2343±0.0003	5.21	FeI;H9;H8;CaII(H+K);D4000;H_δ	0	...	...	...
294	12:52:52.31	-29:27:19.1	1.2455±0.0003	4.13	MgI;CaII(H+K);D4000;CaI	0	...	...	...
304	12:52:54.02	-29:27:18.5	1.2384±0.0003	3.96	MgI;H6;CaII(H+K);D4000;CaI	0	...	...	...
309	12:52:50.66	-29:27:18.1	1.2312±0.0002	3.92	[OII];H6;CaII(H+K);H_δ;CaI	1	0.85±0.20	-32±8	0.5±0.2
310	12:52:52.39	-29:27:18.0	1.2342±0.0003	3.69	FeI;HeI;H8;MgI;H6;CaII(H+K);D4000;CaI	0	...	...	...
313	12:52:55.22	-29:27:16.7	1.2455±0.0006	3.58	FeI;H10;MgI;CaII(H+K);D4000;CaI	0	...	...	...
330	12:52:58.67	-29:27:10.4	1.2297±0.0003	4.09	HeI;H9;H8;MgI;H6;CaII(H+K);D4000;H_δ	0	...	...	...
338	12:52:53.93	-29:27:09.9	1.2312±0.0003	4.09	FeI;HeI;H8;MgI;CaII(H+K);D4000;H_δ;CaI	0	...	...	...
339	12:52:55.63	-29:27:09.7	1.2274±0.0002	3.43	[OII];H9;H8	1	1.78±0.20	-34±4	1.0±0.4
345	12:52:46.98	-29:27:07.7	1.2379±0.0004	1.00	H10;MgI;[OII];D4000	1	5.00±2.85	-40±23	2.9±2.4
366	12:52:48.64	-29:26:59.7	1.2371±0.0001	6.92	[OII]	1	0.08±0.03	-106±36	0.05±0.03
367	12:52:57.27	-29:27:02.3	1.2295±0.0003	6.92	MgI;CaII(H+K);D4000	0	...	...	...
370	12:53:06.41	-29:27:02.4	1.2373±0.0010	1.24	FeI;H10;MgI;CaII(H+K);D4000	0	...	...	...
407	12:53:00.00	-29:26:09.8	1.2354±0.0002	5.92	H8;MgI;CaII(H+K);D4000;H_δ;CaI	0	...	...	...
419	12:52:41.81	-29:26:45.5	1.2338±0.0001	5.07	[OII]	1	1.71±0.72	-79±33	1.0±0.7
432	12:52:58.19	-29:26:41.6	1.2416±0.0003	4.19	H10;MgI;H6;CaII(H+K);D4000	0	...	...	...
442	12:52:54.52	-29:26:39.8	1.2306±0.0003	3.63	HeI;MgI;H6;CaII(H+K);D4000	0	...	...	...
445	12:52:59.98	-29:26:27.2	1.2400±0.0004	2.94	[OII];H10;MgI;H6;CaII(H+K)	1	1.33±0.27	-30±6	0.8±0.4
515	12:52:55.44	-29:26:12.2	1.2402±0.0001	5.26	HeI;H9;CaII(H+K);[OII]	1	0.91±0.20	-55±12	0.5±0.3
619	12:53:00.23	-29:25:42.7	1.2432±0.0004	1.38	[OII];CaI	1	6.70±1.56	-28±7	3.9±1.9
726	12:53:05.93	-29:26:31.4	1.2273±0.0001	2.79	[OII]	1	3.33±0.65	-130±25	1.9±0.9
2065	12:53:03.30	-29:25:39.3	1.2327±0.0001	15.54	[OII]	1	5.09±0.51	-145±15	3.0±1.0
3159	12:53:00.53	-29:26:51.2	1.2459±0.0003	1.87	[OII];H10;H9;H8	1	0.56±0.14	-30±7	0.3±0.2
6301	12:53:03.50	-29:24:48.6	1.2408±0.0001	7.36	[OII]	1	4.68±1.16	-80±20	2.7±1.4
6306	12:52:48.68	-29:24:59.3	1.2380±0.0001	4.45	[OII];H9;H8;MgI;H6;CaII(H+K);H_δ	1	3.20±0.47	-44±6	1.9±0.7
7001	12:52:48.61	-29:30:33.5	1.2368±0.0003	2.49	FeI;HeI;[OII];H9;H8;MgI;H6;CaII(H+K);H_δ	1	0.32±0.09	-25±7	0.2±0.1
7005	12:52:49.93	-29:29:52.3	1.2416±0.0006	3.28	HeI;H10;H8;MgI;H6;CaII(H+K);D4000;H_δ	0	...	...	...
9000	12:53:04.81	-29:26:58.6	1.2438±0.0001	3.20	[OII]	1	0.30±0.01	-100±42	0.2±0.1

[a] Most prominent spectral features: FeI(λ3581), HeI(λ3614), [OII](λ3727), H10(λ3750), H9(λ3770), H8(λ3799), MgI(λ3834), H6(λ3889), the CaII lines (H(λ3934) and K (λ3969)), the decrement at 4000 Å (D4000), H_δ(λ4102) and CaI(λ4227).

TABLE 3
PHOTOMETRIC CATALOG OF SPECTROSCOPICALLY CONFIRMED CLUSTER MEMBERS. THE GALAXY ID IN THE CATALOG IS GIVEN IN THE FIRST COLUMN. MAGNITUDES IN THE AB SYSTEM, CORRECTED FOR GALACTIC REDDENING, AND THE CORRESPONDING ERRORS ARE GIVEN IN THE FOLLOWING COLUMNS: B-, V- AND R-BAND MAGNITUDES WERE OBTAINED WITH VLT/FORS, THE i_{775}- AND z_{850}-BAND MAGNITUDES ARE FROM HST/ACS AND THE NEAR-IR MAGNITUDES ARE FROM THE REPROCESSED VLT/ISAAC DATA. A VALUE OF 99 HAS BEEN ASIGNED TO UNAVAILABLE PHOTOMETRY. THE LAST COLUMN GIVES THE MORPHOLOGICAL T-TYPE (DE VAUCOULEURS ET AL. 1976) OF THE GALAXY AS DETERMINED BY POSTMAN ET AL. (2005).

ID	K_s-Tot	$B-V$	$V-R$	$V-i_{775}$	$R-K_s$	$i_{775}-z_{850}$	J_s-K_s	T-type
137	20.97±0.01	0.60±0.15	1.42±0.08	2.56±0.06	3.72±0.05	0.93±0.02	0.76±0.02	-5
149	21.01±0.01	1.21±0.31	1.68±0.09	2.80±0.08	3.84±0.05	0.90±0.02	0.84±0.02	-2
174	20.98±0.02	0.27±0.10	0.51±0.08	1.37±0.06	3.85±0.06	0.74±0.05	1.25±0.03	8
205	21.03±0.02	1.24±0.26	1.00±0.08	2.20±0.06	3.94±0.06	0.92±0.02	0.80±0.02	-2
206	20.96±0.01	0.52±0.15	1.32±0.08	2.29±0.07	3.78±0.05	0.98±0.03	0.81±0.02	-5
247	19.53±0.01	1.22±0.11	1.26±0.04	2.38±0.03	4.00±0.03	0.99±0.01	0.82±0.01	-5
248	21.68±0.07	0.21±0.08	0.32±0.06	0.96±0.04	2.79±0.06	0.68±0.02	1.00±0.04	8
265	21.02±0.01	0.79±0.15	1.04±0.07	2.15±0.05	3.82±0.05	0.97±0.02	0.85±0.02	-2
282	20.60±0.05	1.02±0.20	1.13±0.08	2.31±0.06	4.28±0.06	0.95±0.01	0.96±0.02	-5
289	19.54±0.04	0.80±0.12	1.45±0.06	2.49±0.04	4.02±0.04	0.96±0.01	0.87±0.01	-5
291	19.22±0.03	1.11±0.11	1.28±0.04	2.34±0.03	4.09±0.03	0.96±0.01	0.91±0.01	-5
294	20.31±0.05	1.08±0.13	1.21±0.06	2.23±0.04	3.97±0.04	0.93±0.01	0.89±0.01	-5
304	19.75±0.11	1.28±0.18	1.45±0.06	2.58±0.05	4.25±0.04	0.97±0.01	0.94±0.02	-5
309	20.97±0.01	0.66±0.09	0.78±0.05	1.47±0.03	3.27±0.04	0.94±0.01	0.87±0.02	-2
310	20.54±0.03	1.08±0.15	0.99±0.06	2.00±0.04	3.99±0.05	0.89±0.01	0.91±0.02	-5
313	20.32±0.01	99	1.52±0.09	2.74±0.07	4.45±0.05	1.02±0.01	1.00±0.01	-5
330	20.85±0.01	0.69±0.16	1.39±0.08	2.55±0.06	3.94±0.05	0.88±0.02	0.86±0.02	-5
338	20.93±0.02	1.18±0.22	1.28±0.08	2.33±0.06	3.92±0.05	0.92±0.01	0.88±0.03	-5
339	20.85±0.01	0.04±0.04	0.20±0.04	0.63±0.02	2.40±0.03	0.64±0.01	0.92±0.02	8
345	20.76±0.01	0.33±0.10	1.07±0.06	2.09±0.04	3.71±0.04	0.99±0.01	0.85±0.02	3
366	22.26±0.08	0.17±0.04	0.14±0.04	0.65±0.02	1.71±0.06	0.52±0.01	0.55±0.06	3
367	20.97±0.04	0.69±0.14	1.05±0.07	2.03±0.05	3.90±0.05	0.95±0.02	0.78±0.02	-2
370	21.33±0.06	99	1.76±0.14	3.07±0.12	4.32±0.11	1.00±0.03	1.00±0.14	-1
407	20.21±0.01	1.13±0.21	1.62±0.07	2.78±0.06	4.29±0.04	0.99±0.02	0.92±0.01	-2
419	21.89±0.08	0.23±0.06	0.32±0.05	0.83±0.04	2.49±0.14	0.69±0.02	0.71±0.19	8
432	21.47±0.01	0.69±0.20	0.91±0.11	2.35±0.08	4.22±0.08	0.81±0.04	0.86±0.03	-5
442	21.26±0.05	0.39±0.17	1.40±0.09	2.59±0.07	3.84±0.06	0.91±0.01	0.77±0.02	-5
445	20.31±0.01	0.44±0.07	0.83±0.05	1.63±0.04	3.51±0.04	0.91±0.02	0.84±0.01	-5
515	22.07±0.11	0.03±0.04	0.24±0.04	0.70±0.02	1.65±0.06	0.46±0.01	0.48±0.09	8
619	21.65±0.02	0.20±0.06	0.27±0.05	0.93±0.04	2.36±0.05	0.70±0.04	0.61±0.04	6
726	23.15±0.16	0.06±0.06	0.26±0.05	0.65±0.04	1.12±0.43	0.33±0.04	0.18±0.52	-
2065	22.78±0.05	0.09±0.04	0.20±0.04	0.58±0.03	1.23±0.07	0.43±0.03	0.34±0.08	8
3159	23.87±0.09	0.55±0.05	0.23±0.04	0.52±0.03	0.62±0.12	0.14±0.03	0.07±0.13	0
6301	99	0.00±0.04	0.22±0.04	0.48±0.04	99	0.43±0.04	99	0
6306	99	0.29±0.07	0.63±0.05	1.22±0.04	99	0.85±0.03	99	1
7001	99	99	99	99	99	99	99	-
7005	99	1.06±0.21	1.16±0.08	2.14±0.07	99	0.88±0.04	99	-1
9000	99	0.23±0.11	0.36±0.09	0.87±0.08	99	0.61±0.08	99	8

TABLE 4
X-RAY POINT SOURCES, IN THE CLUSTER FIELD OF VIEW, DETECTED BY CHANDRA. RIGHT ASCENSION (R.A.) AND DECLINATION (DEC.) J2000 COORDINATES ARE IN (H:M:S) AND (^{o}:$'$:$''$), RESPECTIVELY. THE SPECTROSCOPIC REDSHIFT OF THE SOURCE IS GIVEN IN COLUMN 4. THE FLUX, IN UNITS OF $erg\ s^{-1}\ \ cm{-}2$, IN THE [0.5-2] AND [2-10] KEV BANDS OF CHANDRA ARE GIVEN IN COLUMNS 5 AND 6, RESPECTIVELY. THE REST FRAME X-RAY LUMINOSITY, IN UNITS OF $erg\ s^{-1}$, IN THE [0.5-2] AND [2-10] KEV BANDS ARE GIVEN IN COLUMNS 7 AND 8, RESPECTIVELY.

ID	R.A.	DEC.	z	f_x[0.5-2]	f_x[2-10]	L_x[0.5-2]	L_x[2-10]
425	12:52:58.14	-29:26:49.3	0.4694	5.1E-16	2.3E-15	3.3E+41	1.5E+42
69	12:53:03.60	-29:28:28.4	0.4762	1.1E-15	6.8E-15	7.3E+41	4.6E+42
7006	12:52:56.96	-29:30:10.9	0.7313	1.4E-15	3.3E-15	2.4E+42	5.7E+42
661	12:53:01.50	-29:25:38.8	0.7439	1.4E-14	2.6E-14	2.6E+43	4.8E+43
217	12:53:03.64	-29:27:42.1	0.7533	3.3E-16	9.0E-16	6.1E+41	1.7E+42
3043	12:53:00.39	-29:29:17.2	0.8220	2.9E-15	1.1E-14	6.5E+42	2.4E+43
355	12:52:42.48	-29:27:03.0	0.8440	4.2E-16	9.5E-16	9.9E+41	2.3E+42
612	12:52:59.93	-29:25:18.4	1.1764	1.1E-15	4.5E-15	5.5E+42	2.2E+43
174	12:52:49.82	-29:27:54.9	1.2382	1.9E-16	1.6E-15	1.1E+42	8.6E+42
321	12:52:40.36	-29:27:14.1	1.3479	1.6E-16	5.6E-15	1.1E+42	3.6E+43
374	12:52:50.08	-29:27:00.7	1.5198	1.6E-15	3.5E-15	1.4E+43	3.0E+43

TABLE 5
SPECTROSCOPIC DATA OF FIELD GALAXIES. THE OBJECT ID IN THE CATALOG IS GIVEN IN THE FIRST COLUMN. RIGHT ASCENSION (R.A.) AND DECLINATION (DEC.) J2000 COORDINATES ARE GIVEN IN THE SECOND AND THIRD COLUMN RESPECTIVELY. RAS ARE GIVEN IN (H:M:S) AND DECS IN (^{o}:$'$:$''$). COLUMNS 4 AND 5 SHOW THE REDSHIFT AND THE FORMAL ERROR IN REDSHIFT AS OBTAINED FROM THE CROSS-CORRELATION. THE TABLE HAS BEEN ORDERED FROM LOW TO HIGH REDSHIFT. A PROMINENT "STRUCTURE" IN REDSHIFT CAN BE SEEN AT REDSHIFT $z \sim 0.74$ (SEE SECTION §3.1). OBJECTS BELONGING TO THIS "STRUCTURE" ARE MARKED WITH A STAR.

ID	R.A.	DEC.	z	$Err.\ z$
128	12:53:05.23	-29:28:08.2	0.0811	0.0008
723	12:52:44.53	-29:29:10.0	0.0844	0.0012
413	12:52:48.10	-29:26:50.7	0.0906	0.0011
29	12:52:55.14	-29:28:49.5	0.1092	0.0009
787	12:52:49.99	-29:30:00.7	0.1637	0.0004
718	12:53:00.43	-29:24:35.3	0.1784	0.0010
624	12:52:59.72	-29:25:41.6	0.1798	0.0008
179	12:52:59.88	-29:27:50.4	0.1952	0.0007
502	12:52:40.33	-29:26:20.2	0.2461	0.0005
273	12:52:44.34	-29:27:29.6	0.2471	0.0001
219	12:52:44.27	-29:27:40.1	0.2476	0.0005
3096	12:52:45.39	-29:29:27.5	0.2604	0.0002
371	12:52:47.38	-29:27:02.5	0.2613	0.0008
31	12:52:48.07	-29:28:42.2	0.2615	0.0007
535	12:52:51.52	-29:25:58.1	0.2693	0.0010
792	12:53:09.19	-29:27:36.3	0.2781	0.0006
746	12:53:05.03	-29:26:15.2	0.2788	0.0007
162	12:52:44.63	-29:28:00.3	0.2882	0.0007
594	12:52:48.67	-29:25:10.9	0.2979	0.0004
449	12:52:49.21	-29:26:39.1	0.3060	0.0001
72	12:53:06.18	-29:28:28.6	0.3233	0.0005
241	12:52:57.98	-29:27:36.2	0.3305	0.0005
250	12:52:46.42	-29:27:28.3	0.3452	0.0014
4052	12:53:06.44	-29:26:50.1	0.3483	0.0004
361	12:53:00.97	-29:27:04.0	0.3492	0.0001
341	12:52:54.73	-29:27:10.5	0.3515	0.0007
27	12:52:50.15	-29:28:53.9	0.3749	0.0005
467	12:52:53.07	-29:26:31.4	0.3760	0.0006
325	12:52:40.53	-29:27:14.0	0.3962	0.0001
181	12:53:09.80	-29:27:51.6	0.4013	0.0004
180	12:52:56.32	-29:27:51.4	0.4050	0.0007
365	12:52:58.61	-29:27:03.0	0.4058	0.0004
790	12:52:59.31	-29:29:57.0	0.4296	0.0006
736	12:53:03.33	-29:24:17.1	0.4346	0.0006
730	12:52:44.87	-29:27:00.7	0.4602	0.0002
4010	12:53:10.59	-29:26:02.9	0.4687	0.0002
552	12:53:08.59	-29:25:58.4	0.4687	0.0006
629	12:52:56.07	-29:25:25.2	0.4689	0.0004
492	12:52:54.71	-29:26:23.1	0.4690	0.0004
269	12:52:56.13	-29:27:22.1	0.4692	0.0004
425	12:52:58.14	-29:26:49.3	0.4694	0.0006
69	12:53:03.60	-29:28:28.4	0.4762	0.0005
22	12:52:47.45	-29:28:56.7	0.4827	0.0004
714	12:53:09.89	-29:26:37.1	0.4861	0.0006
571	12:52:44.92	-29:25:47.2	0.4941	0.0007
153	12:52:47.06	-29:28:03.2	0.5110	0.0002
546	12:52:51.31	-29:25:58.1	0.5125	0.0009
656	12:52:48.90	-29:25:37.7	0.5375	0.0004
649	12:52:49.03	-29:25:37.4	0.5390	0.0004
469	12:52:45.78	-29:26:30.9	0.5408	0.0005
717	12:53:03.78	-29:24:34.9	0.5530	0.0005
788	12:52:52.25	-29:30:14.1	0.5573	0.0005
71	12:53:08.30	-29:28:30.3	0.5616	0.0002
756	12:52:40.61	-29:27:10.3	0.5952	0.0006
377	12:53:00.10	-29:27:01.2	0.6678	0.0003
737	12:52:53.31	-29:24:43.3	0.6700	0.0005
127	12:52:59.38	-29:28:11.5	0.6701	0.0002
1	12:52:45.89	-29:29:04.0	0.6705	0.0005
222	12:52:53.09	-29:27:39.6	0.6719	0.0002
193	12:52:51.10	-29:27:49.0	0.6721	0.0002
528	12:52:57.67	-29:26:05.2	0.6723	0.0008
398	12:52:54.57	-29:26:48.5	0.6733	0.0003
405	12:53:00.24	-29:26:47.8	0.6744	0.0006
720	12:52:48.01	-29:29:10.0	0.6802	0.0006
441*	12:52:46.03	-29:26:40.1	0.7201	0.0004
175*	12:52:58.98	-29:27:56.3	0.7217	0.0010
7006*	12:52:56.96	-29:30:10.9	0.7313	0.0002
508*	12:52:54.75	-29:26:14.7	0.7331	0.0003
537*	12:52:54.50	-29:26:05.4	0.7333	0.0003
465*	12:52:52.88	-29:26:31.9	0.7336	0.0003
495*	12:52:56.38	-29:26:22.3	0.7336	0.0002

TABLE 6
Color-magnitude fitting. Only spectroscopic members without [OII] emission are used to fit the CM relation (solid red line in Fig. 14).

Fit #	Slope	Color at $K_s = 18.5$	Intrinsic Scatter	Comments
1	-0.048 ± 0.015	1.85	0.070	(Lidman et al. 2004)
2	-0.042 ± 0.022	1.87	0.061	This work

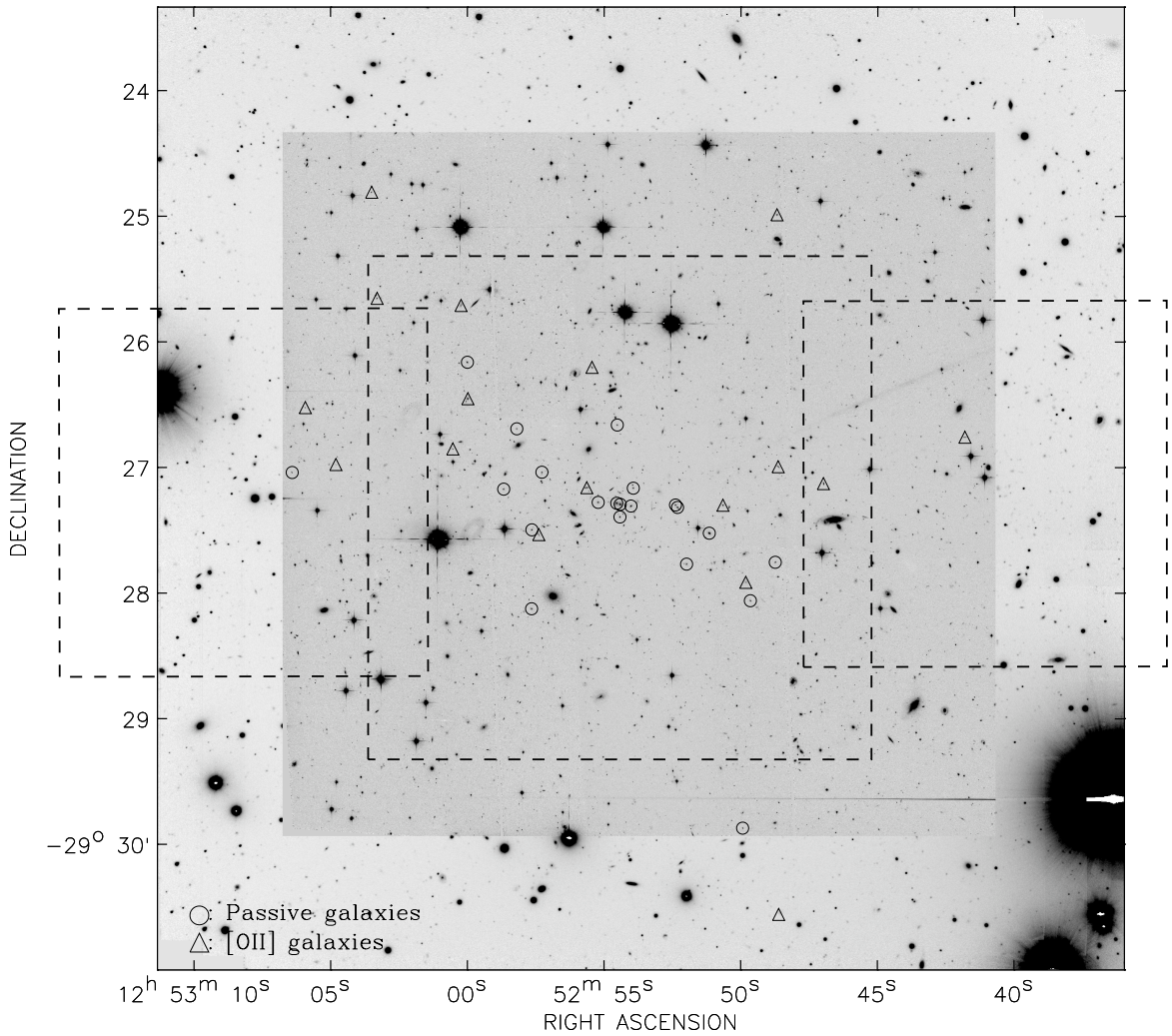

FIG. 1.— Projected distribution of cluster members on the sky (North is up and East is to the left). Circles are non star-forming galaxies and triangles are galaxies showing [OII](λ3727) in their spectrum. The background image corresponds to the FORS2 coverage in BVR, the darker image corresponds to the z_{850} ACS data and the dashed lines show the field covered by ISAAC. The overall shape of the distribution of confirmed galaxy members is clearly elongated in the East-West direction.

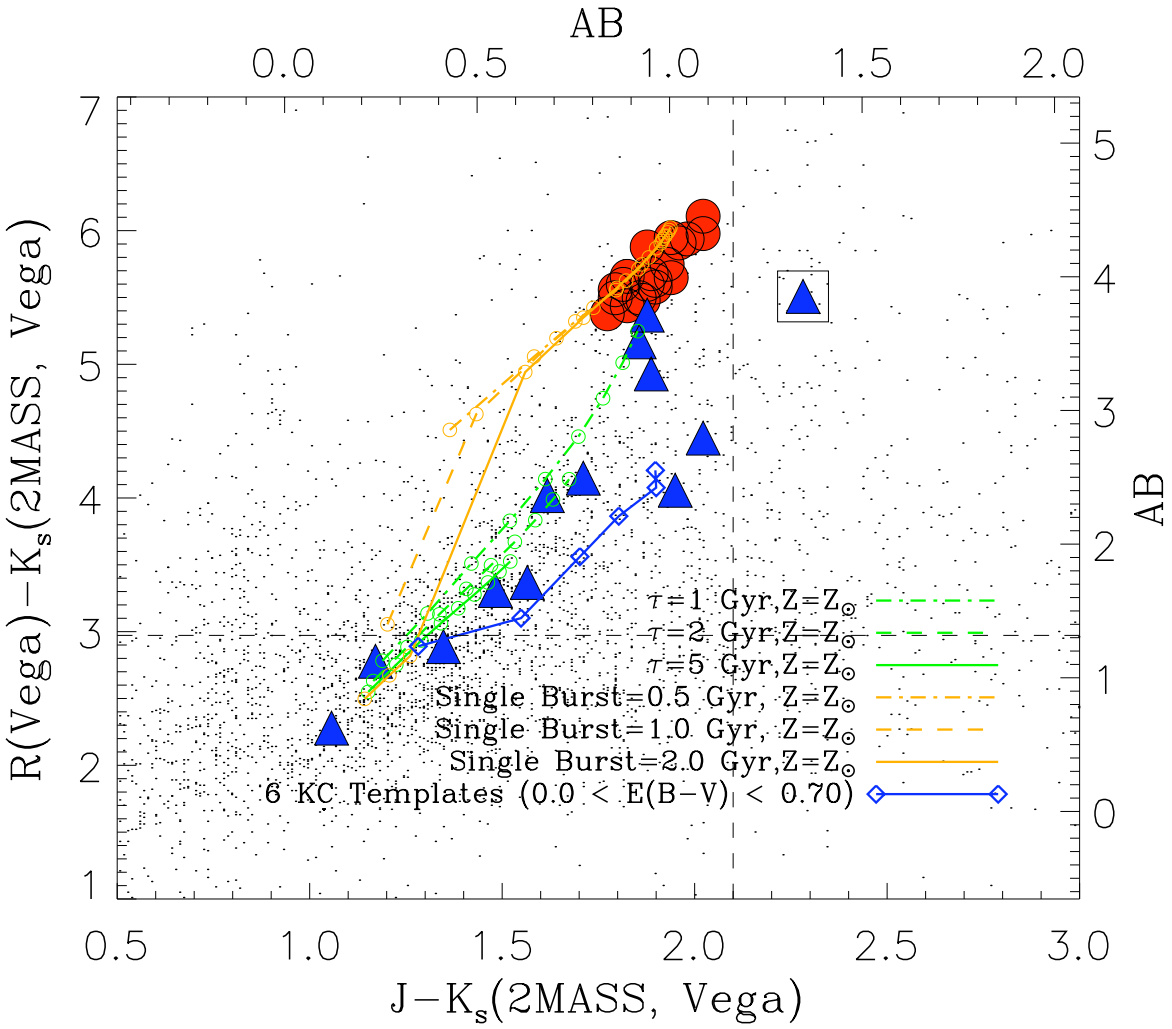

FIG. 2.— Color-color selection of the spectroscopic sample. Galaxies with $K_s < 21$, $J-K_s < 2.1$ and $R-K_s > 3$ were targeted. The different color tracks correspond to different evolutionary models, computed with Bruzual & Charlot's code, and to 6 KC templates with different $E(B-V)$ color excess (see text for details). The square indicates the spectroscopically confirmed cluster AGN (ID=174). Filled circles indicate passive members while filled triangles indicate emission line [OII](λ3727) members. Small dots are objects in our full photometric catalog.

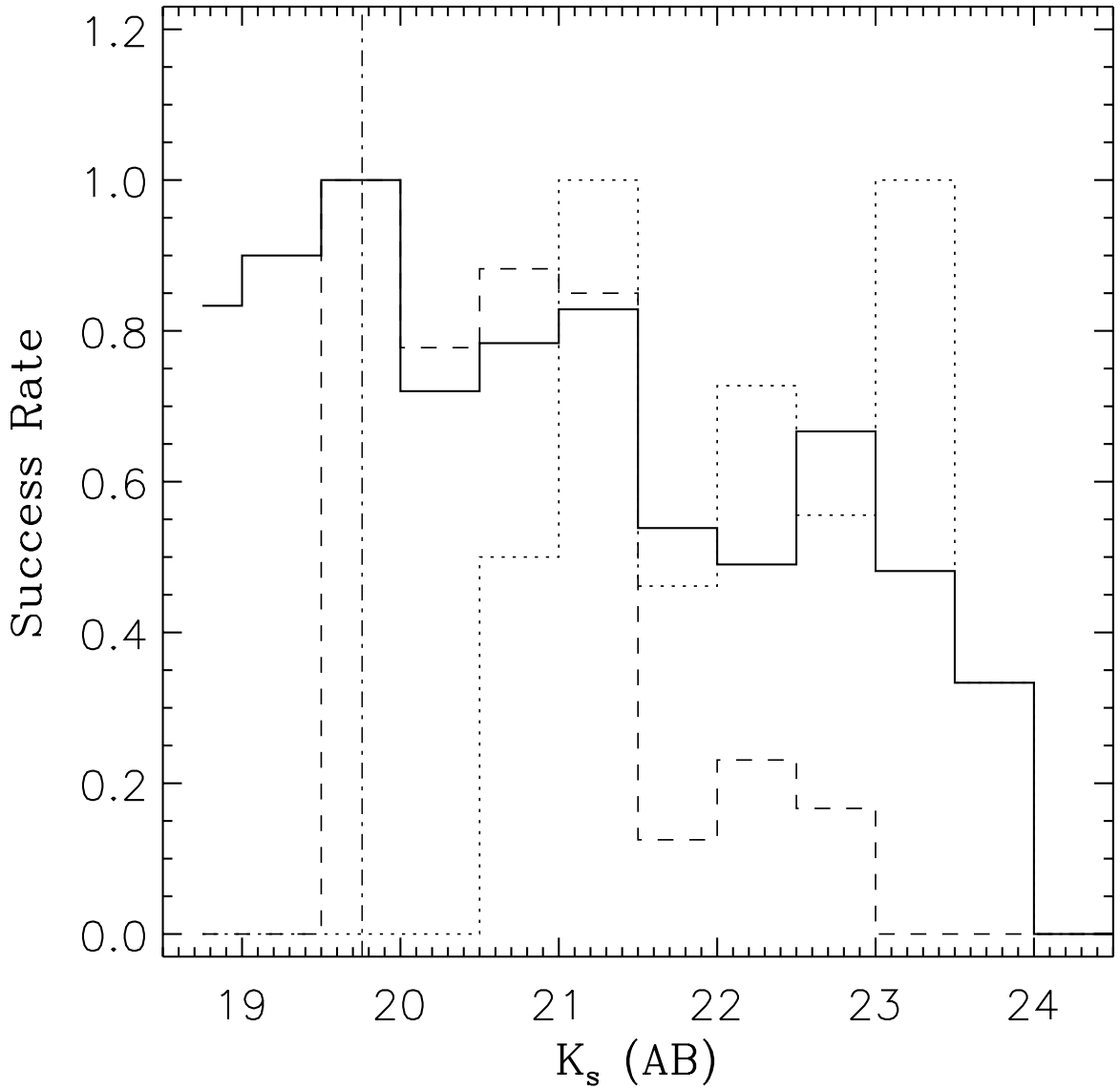

FIG. 3.— The success rate of the spectroscopic survey, defined as the ratio of the number of objects with spectroscopic redshifts to the number of objects that were targeted for spectroscopy as a function of K_s magnitude, is shown by the solid histogram. The data were binned in $\Delta K_s = 0.5$ mag intervals. This distribution includes only objects for which an estimate of their K_s magnitude is available. The dashed histogram shows the success rate only considering our main color-color selection: $R - K_s > 3$ and $J - K_s < 2.1$, while the dotted histogram shows the same ratio only considering the secondary color-color selection: $I - z_{850} > 0.4$, $I - z_{850} < 0.85$, $V - I > 0.2$, $V - I < 1.2$, $V - I > [2.4\,(I - z_{850}) - 1.12]$ and $V - I < [7.0\,(I - z_{850}) - 2.3]$. This secondary selection was intended to target Balmer-Break and star-forming galaxies. The dot-dashed line corresponds to the K_s magnitude of the brightest cluster member.

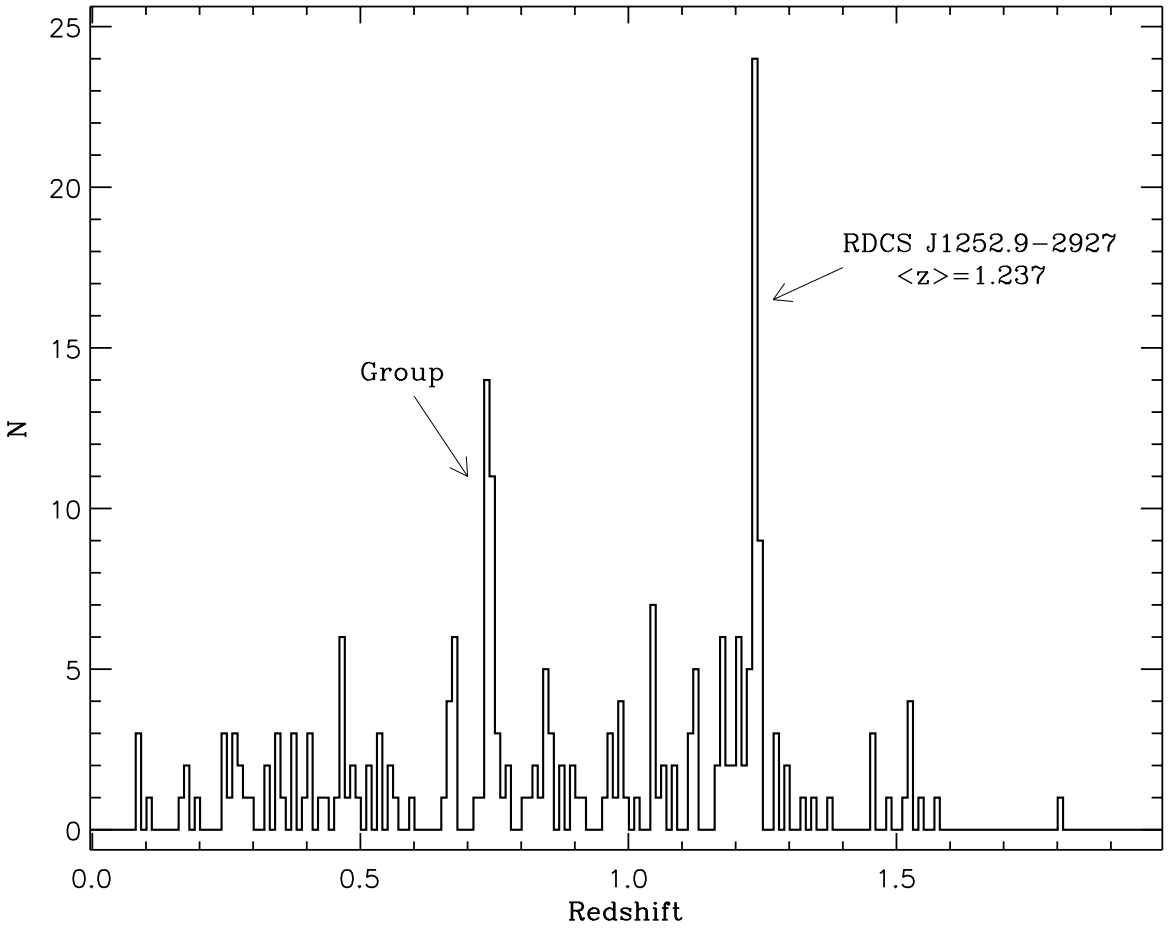

FIG. 4.— Redshift distribution of the 226 galaxies with a secure estimate of their redshifts. The bin size of the histogram is $\Delta z = 0.01$. RDCS J1252.9-2927 appears as the strongest peak of the distribution with a median redshift of $z = 1.2373$ (3σ-clipped mean of $\bar{z} = 1.2371 \pm 0.0009$), with 38 galaxies within the range $1.22 < z < 1.25$. A number of smaller peaks in the survey distribution are observed along the line of sight to the cluster, where a significant peak at $z \simeq 0.74$ is detected. This "group", composed of 31 galaxies, has a 3σ-clipped mean redshift of $\bar{z} = 0.7429 \pm 0.0024$.

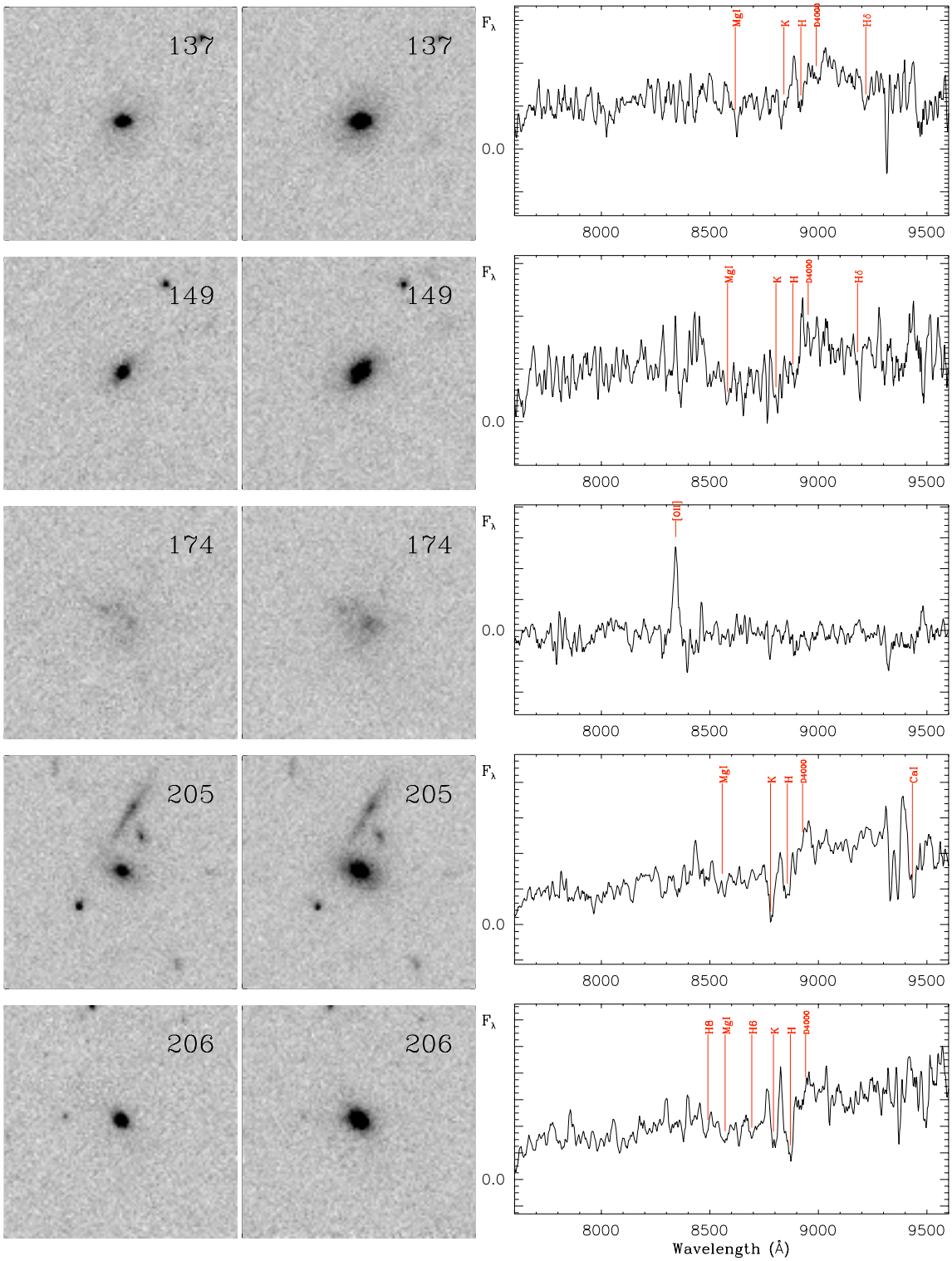

FIG. 5.— The sample of 37 spectroscopic cluster members in the ACS field of view. The left and middle panels show ACS cutouts in the i_{775}- and z_{850}-band respectively. Each cutout is 5″ a side and the number corresponds to the galaxy ID. The right panel shows the corresponding VLT/FORS spectrum. Most prominent spectral features identified in each spectrum in the wavelength range shown here are marked (see footnote in table 2 for the list of spectral features). The spectra have been smoothed by 5 pixels (1 pix $\sim$ 2.5Å).

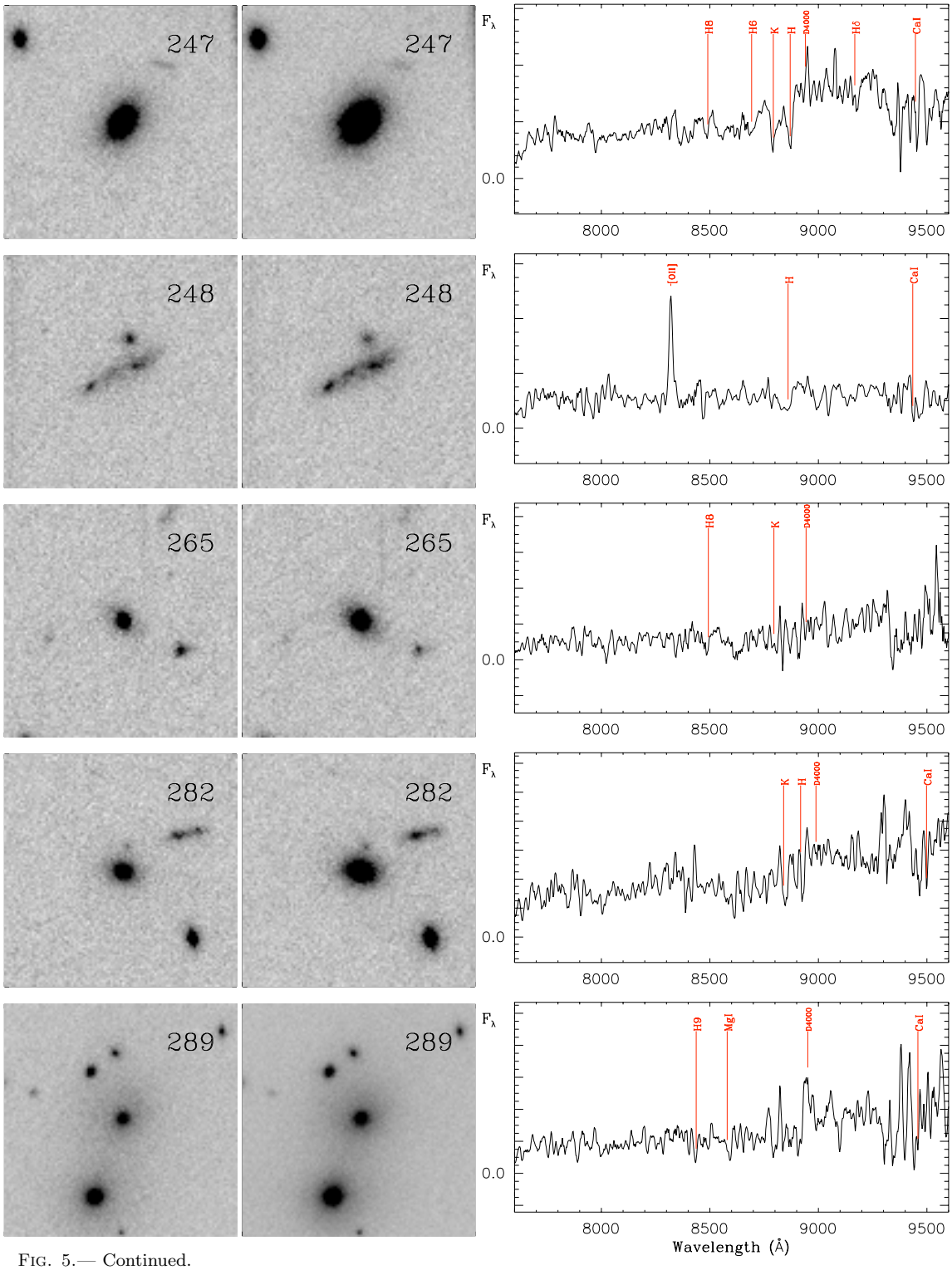

FIG. 5.— Continued.

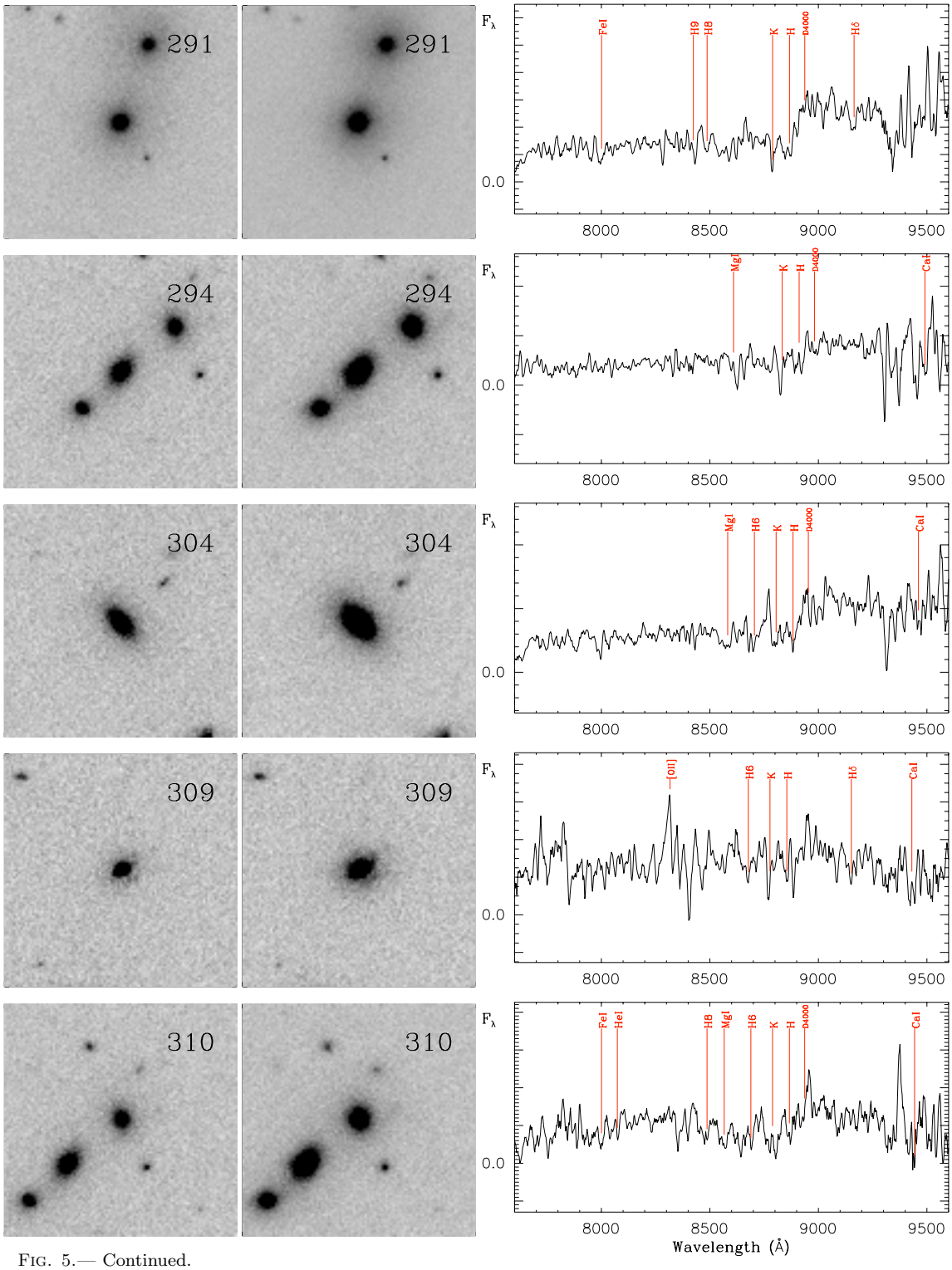

FIG. 5.— Continued.

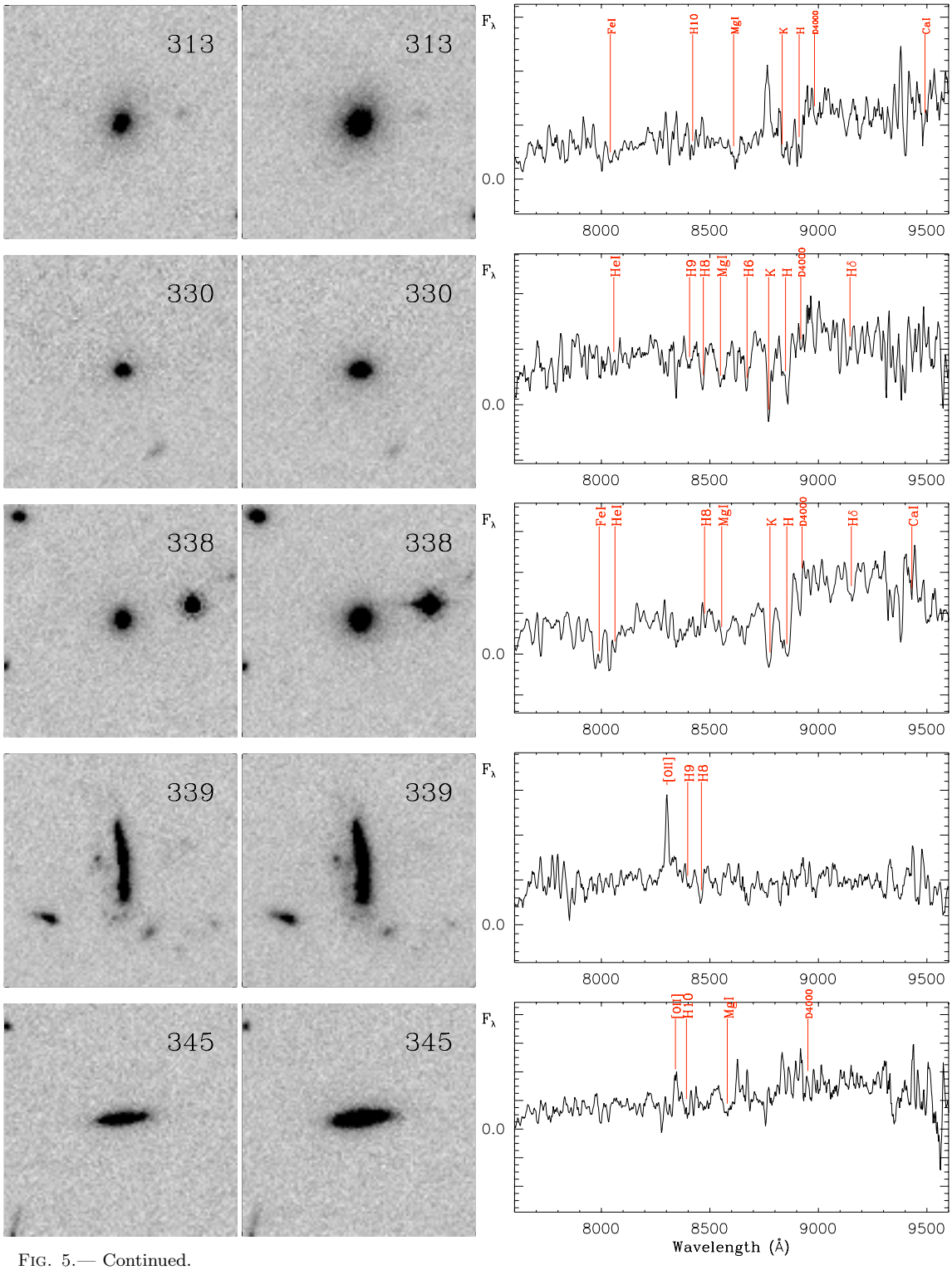

FIG. 5.— Continued.

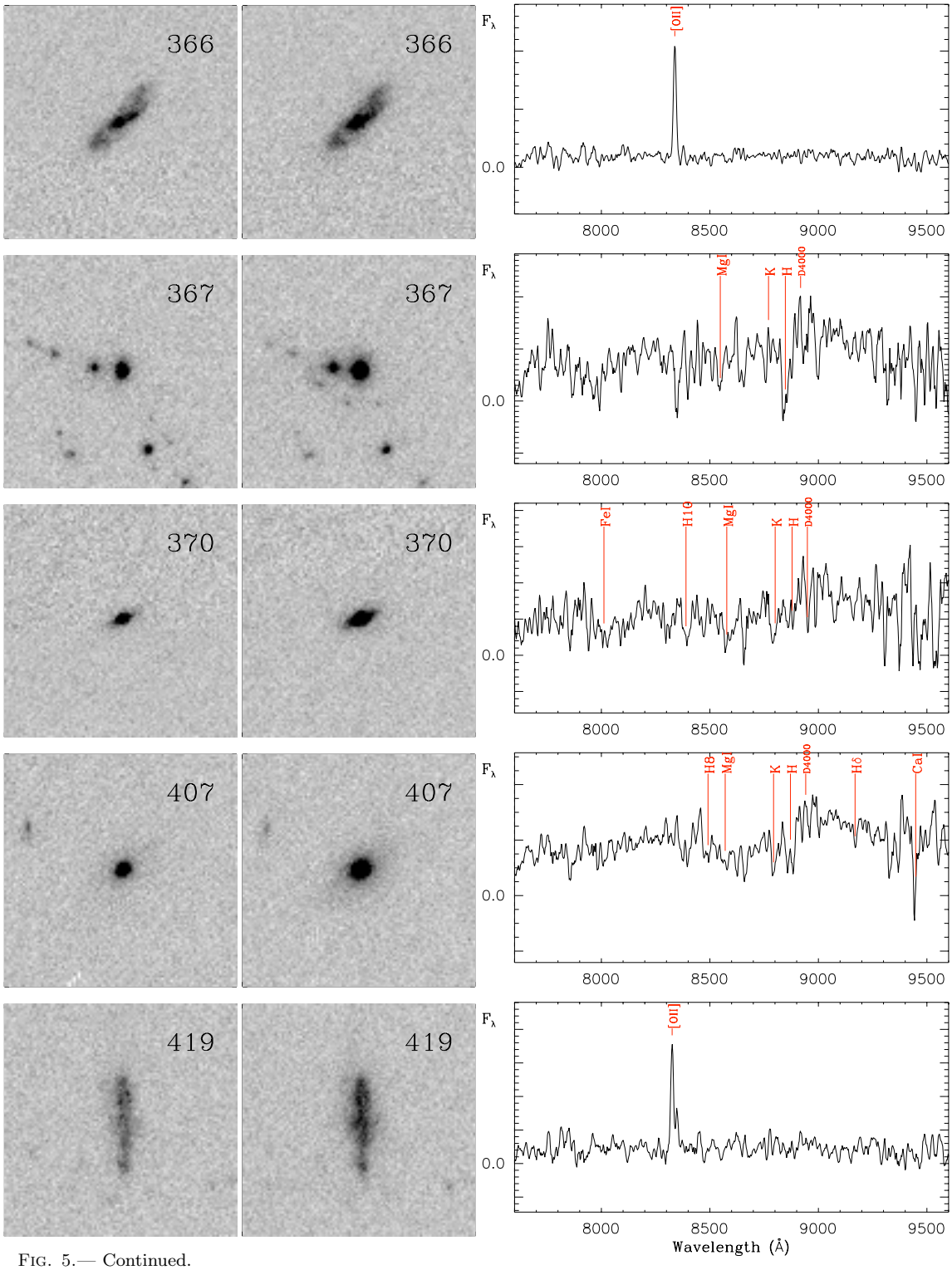

FIG. 5.— Continued.

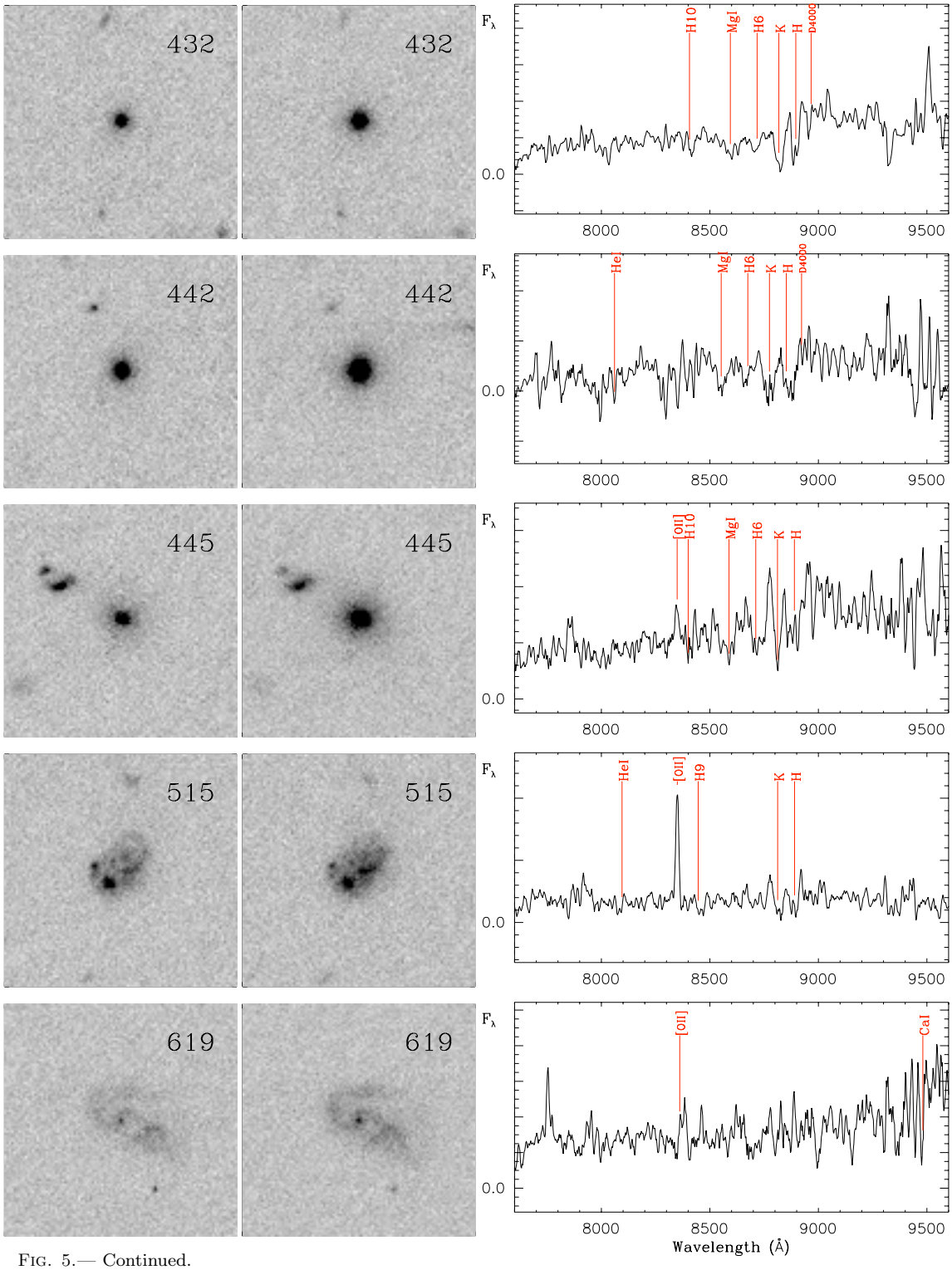

FIG. 5.— Continued.

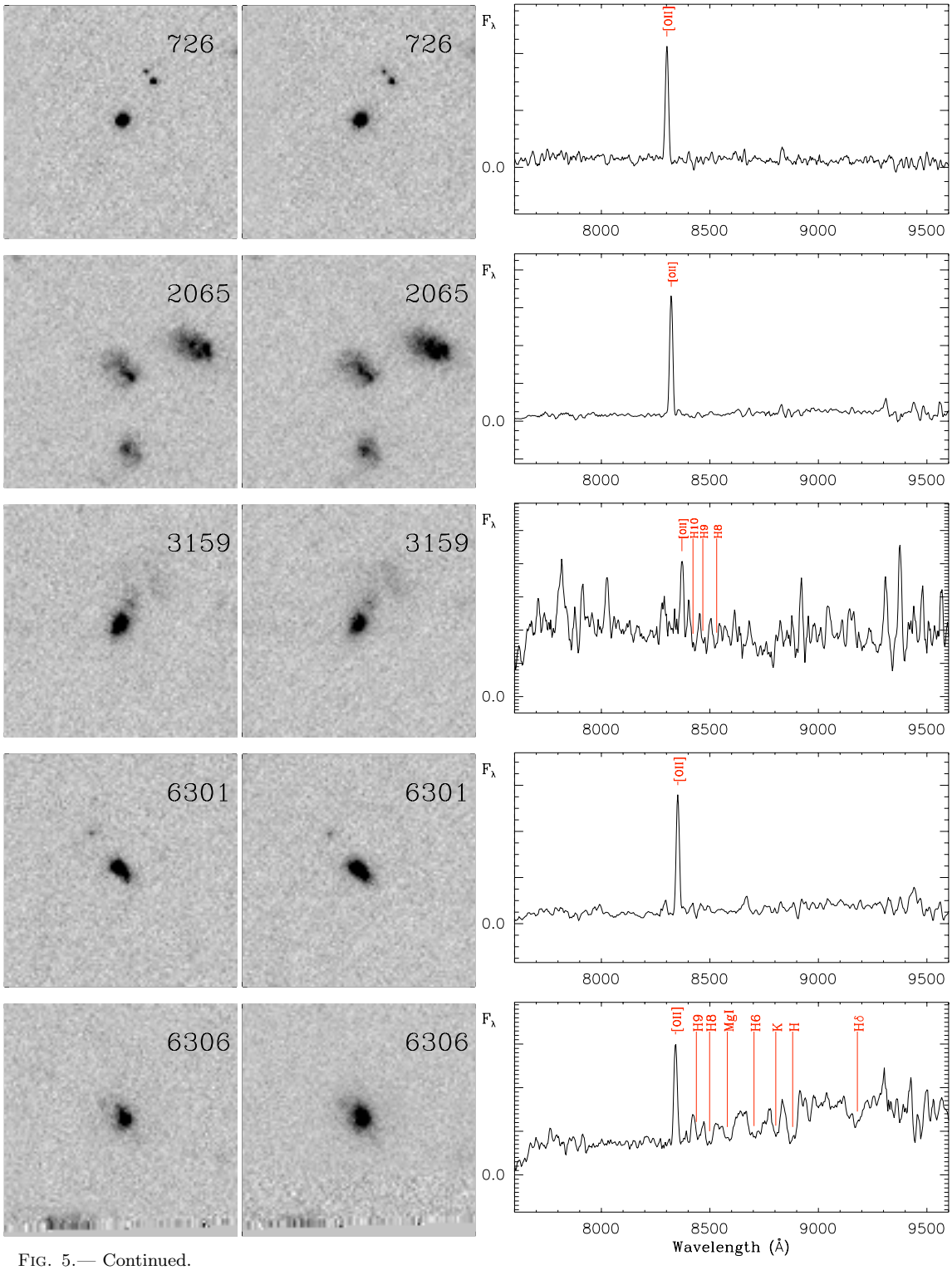

FIG. 5.— Continued.

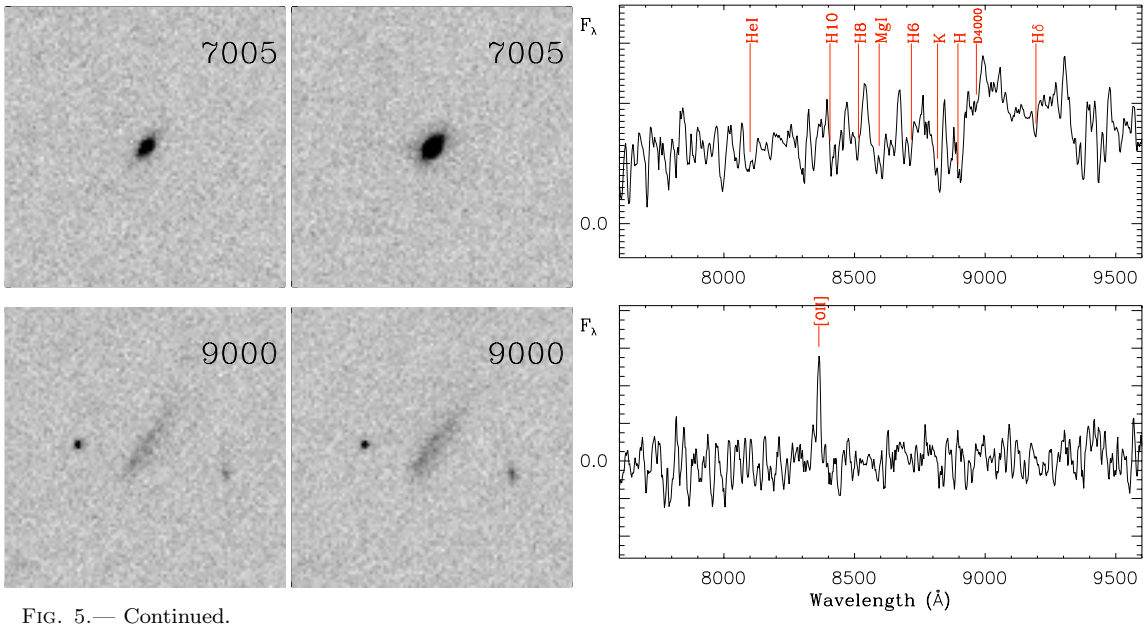

FIG. 5.— Continued.

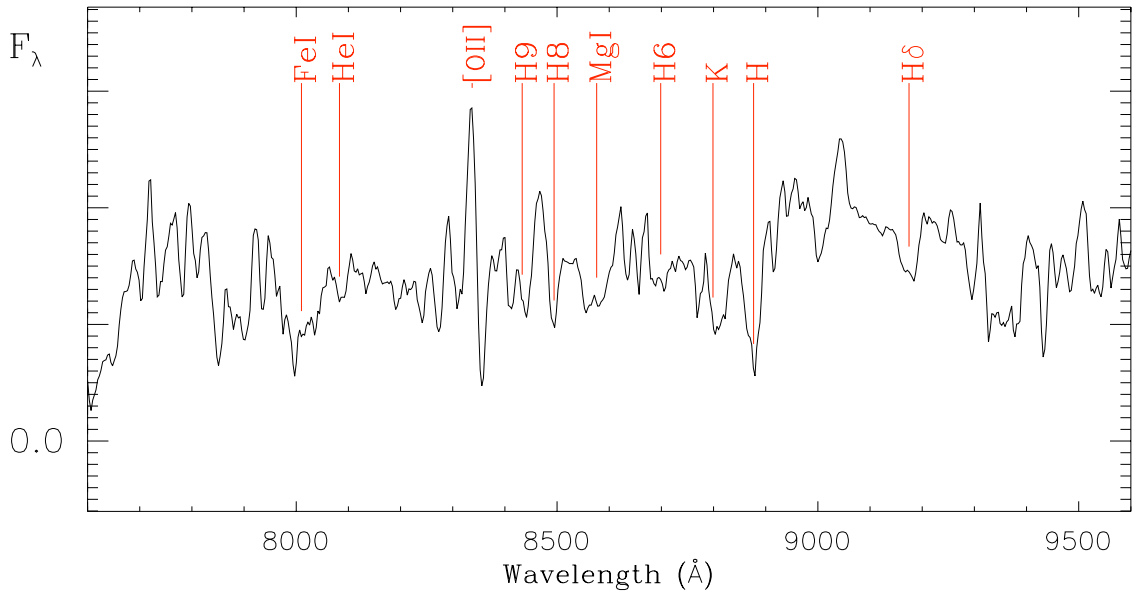

FIG. 6.— VLT/FORS spectrum of cluster member ID=7001. This object was discovered outside the ACS field of view. Most prominent spectral features in the wavelength range shown here are marked (see footnote in table 2 for the list of spectral features). The spectrum has been smoothed by 5 pixels (1 pix $\sim$ 3.1Å).

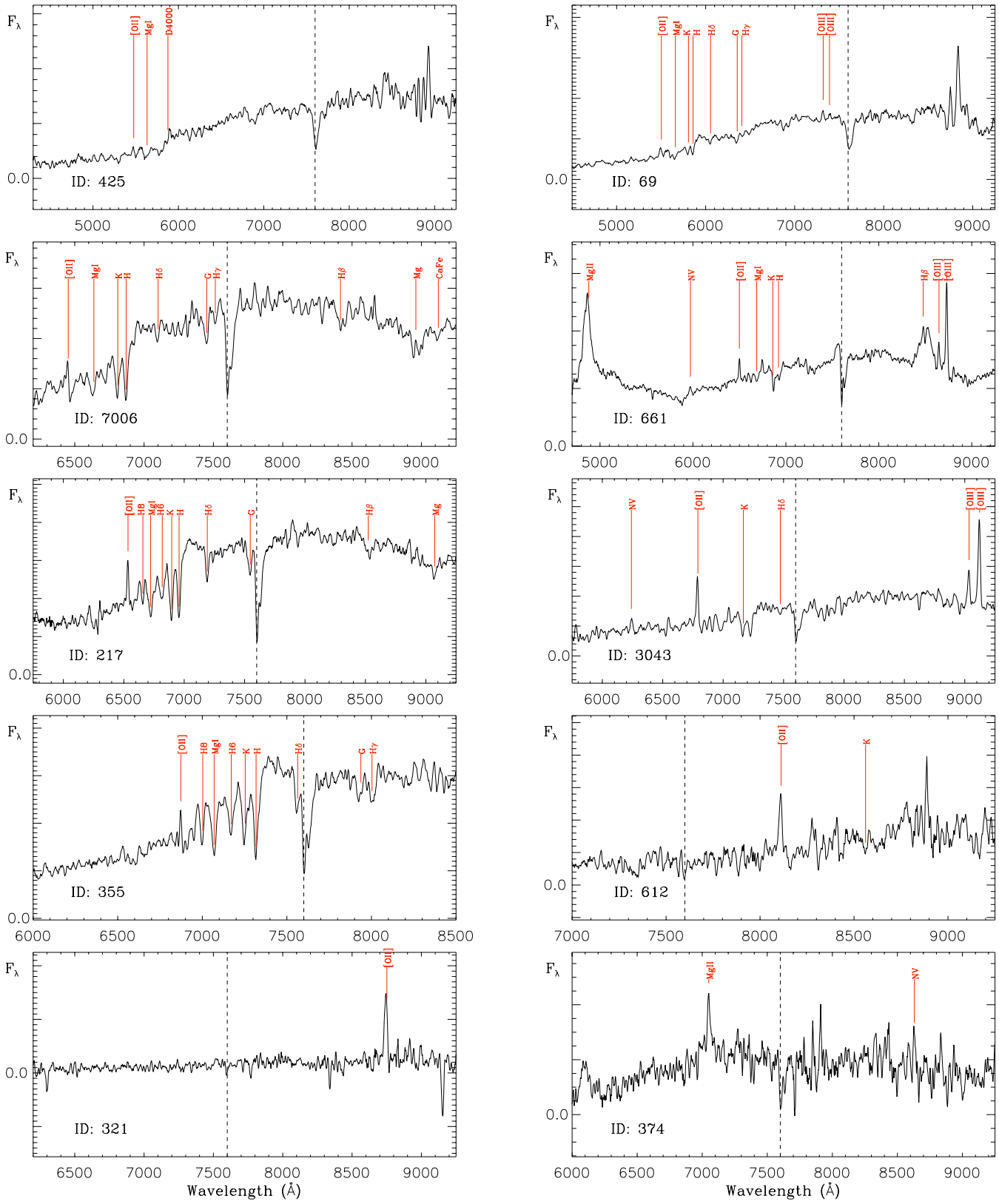

FIG. 7.— Optical spectrum of spectroscopically identified non cluster member AGN in the field of view of RDCS J1252.9-2927. Most prominent spectral features are indicated: MgII(λ2798), NV, [OII](λ3727), H8, MgI(λ3834), H6, CaII(H),CaII(K), D4000, Hδ, G, Hγ, Hβ, [OIII], Mg and CaFe. The vertical dashed line indicates the telluric A-band feature at 7600Å. The spectra have been smoothed by 5 pixels (1 pix $\sim$ 3.0Å).

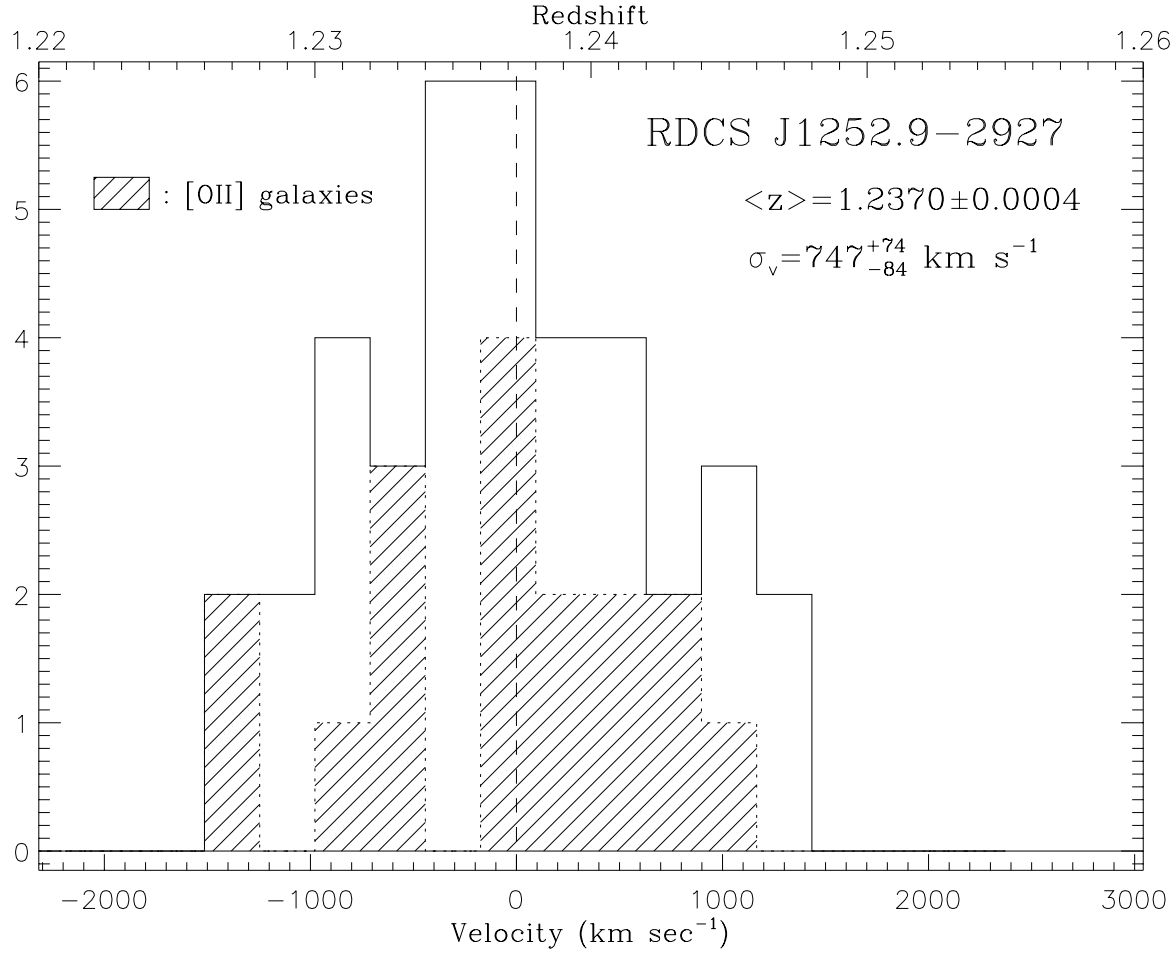

FIG. 8.— Redshift and velocity distribution of cluster members. The histogram bin size has been set to $\Delta z = 0.002$. Assuming a Gaussian distribution in velocity of the spectroscopic members, the median redshift of the distribution is $z = 1.2373$ (vertical dashed line) and the global velocity dispersion is 747^{+74}_{-84} km s^{-1}, based on the biweight estimator (Beers et al. 1990). The hatched area indicates the distribution of star-forming ([OII](λ3727)) members. The mean redshift and velocity dispersion of passive and [OII] cluster members are consistent with each other and with the overall cluster values within the uncertainties.

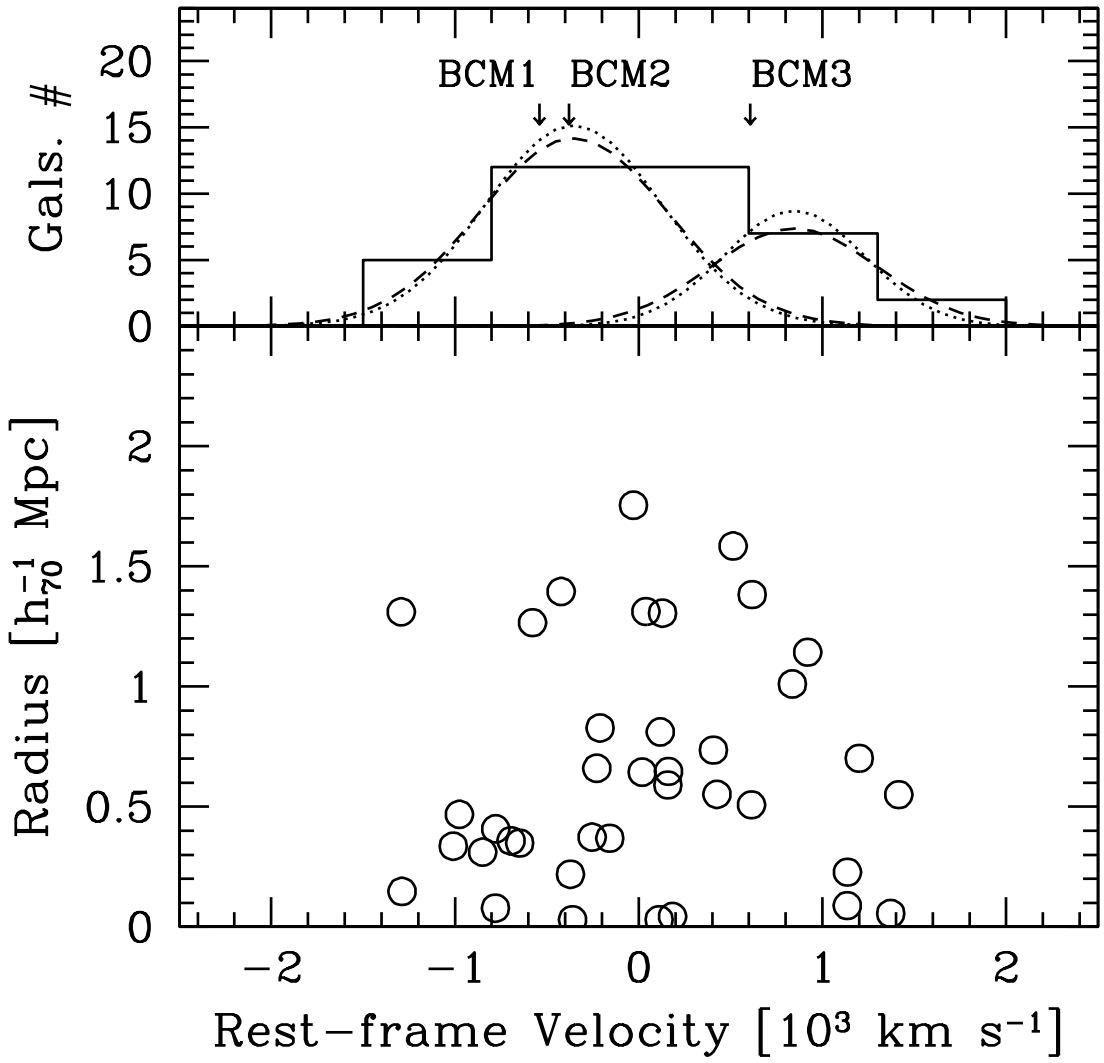

FIG. 9.— *Lower panel*: rest-frame velocity vs. projected distance from the cluster center (taken as the X-ray center; Rosati et al. 2004) of the 38 galaxies assigned to the cluster. *Upper panel*: velocity histogram of the 38 galaxies assigned to the cluster (solid). Velocities of the three brightest galaxies are pointed out. The two dashed Gaussians correspond to the 3D KMM partition (KMM1 and KMM2 from left to right) while the dotted Gaussians indicate the two groups detected by the WGAP procedure (WGAP1 and WGAP2 from left to right). The 3D diagnostics is, in general, the most sensitive indicator of the presence of substructure. Therefore, we consider the KMM groups as a robust characterization of the cluster substructure.

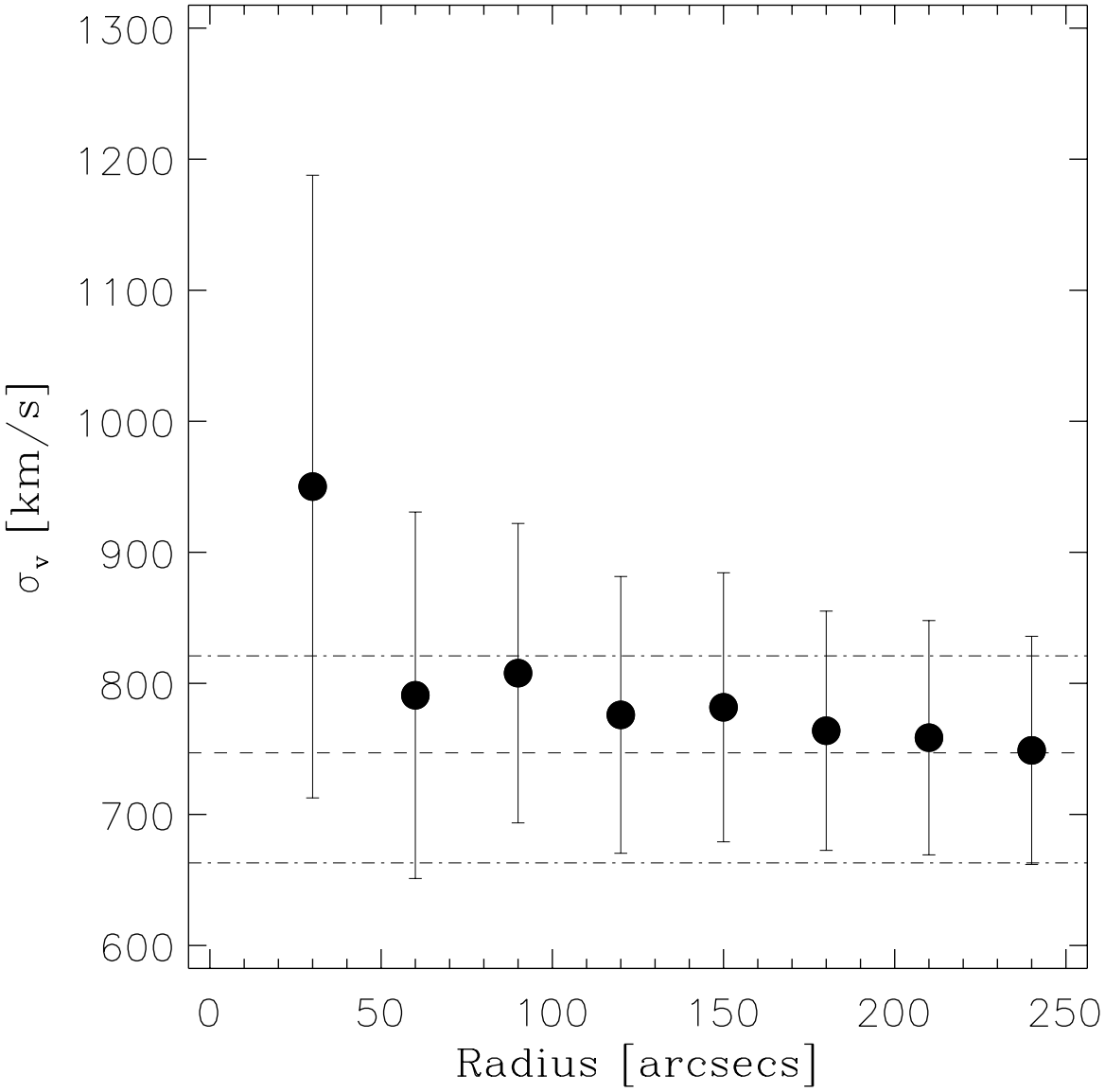

FIG. 10.— Rest-frame integrated velocity dispersion profile of RDCS J1252.9-2927. At the cluster redshift ($z = 1.237$), 0.5 arcmin corresponds to 250 kpc. The dashed line indicates the overall rest-frame cluster velocity dispersion ($\sigma_v = 747^{+74}_{-84}\ km\ s^{-1}$) obtained from the biweight estimator (Beers et al. 1990) with the corresponding error bars as dot-dashed lines. The solid dots are the integrated rest-frame velocity dispersion calculated at the given radius by using all the galaxies within that radius (Girardi et al. 1996). These values were obtained by using Tukey's biweight estimator (Press et al. 1992), correcting for velocity errors (Danese et al. 1980). Error bars are obtained from the fractional uncertainty in estimate for σ_v (Taylor 1997). The cluster velocity dispersion values are observed to be robust at radii larger than 170 arcsec (1.4 Mpc), i.e., at these radii, any velocity anisotropy of cluster galaxies does not affect the value of $\sigma_v(< R)$.

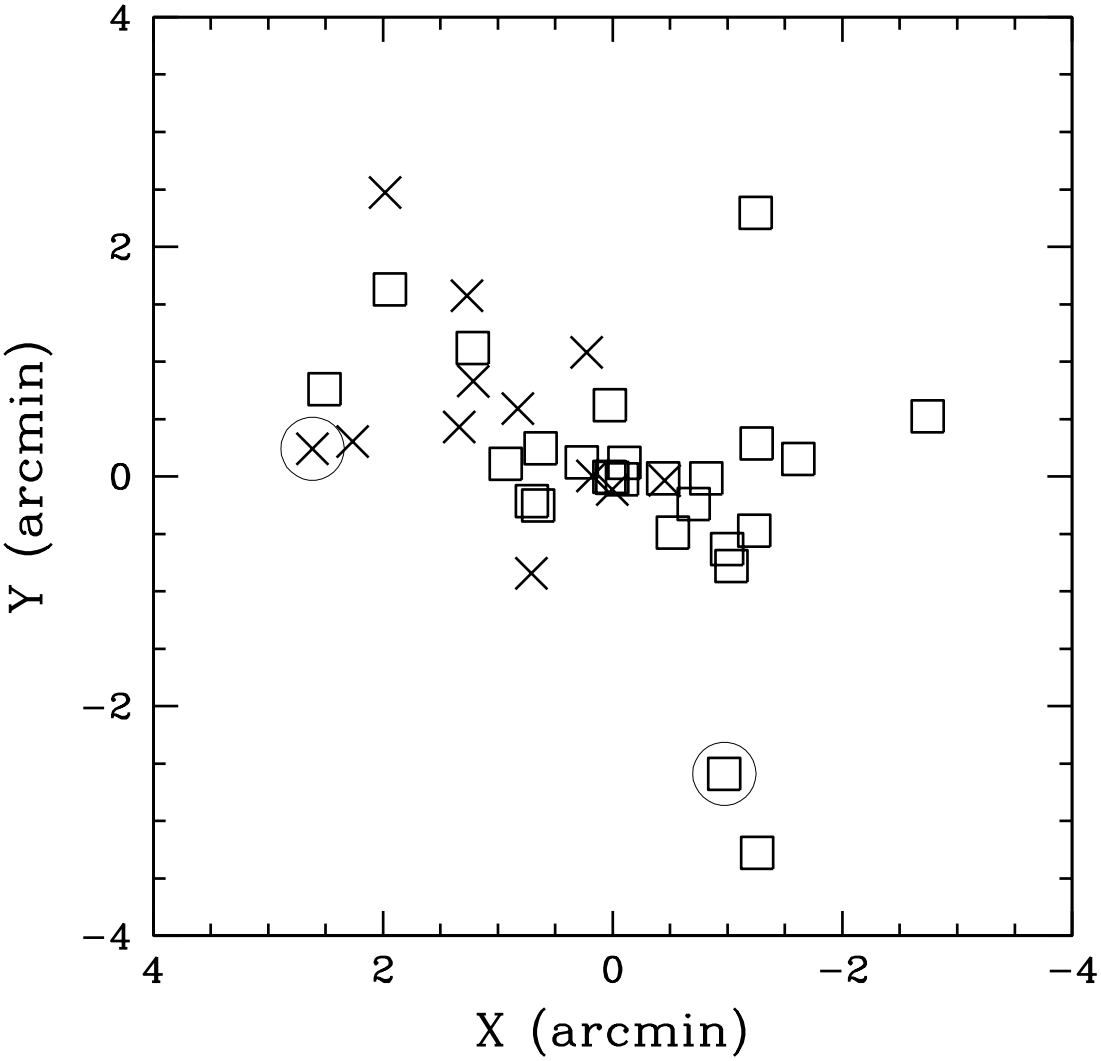

FIG. 11.— Projected distribution on the sky of the 38 member galaxies: open squares and crosses indicate galaxies assigned to KMM1 and KMM2 groups, respectively. The two large open circles indicate the two galaxies which gives the difference in membership with WGAP1 and WGAP2 groups. The plot is centered on the X–ray cluster center. KMM1 is mostly located to the West of the cluster center while KMM2 is mostly to the East, and the degree of merging observed from galaxy positions shows that these two groups have already started virialization in the main cluster potential.

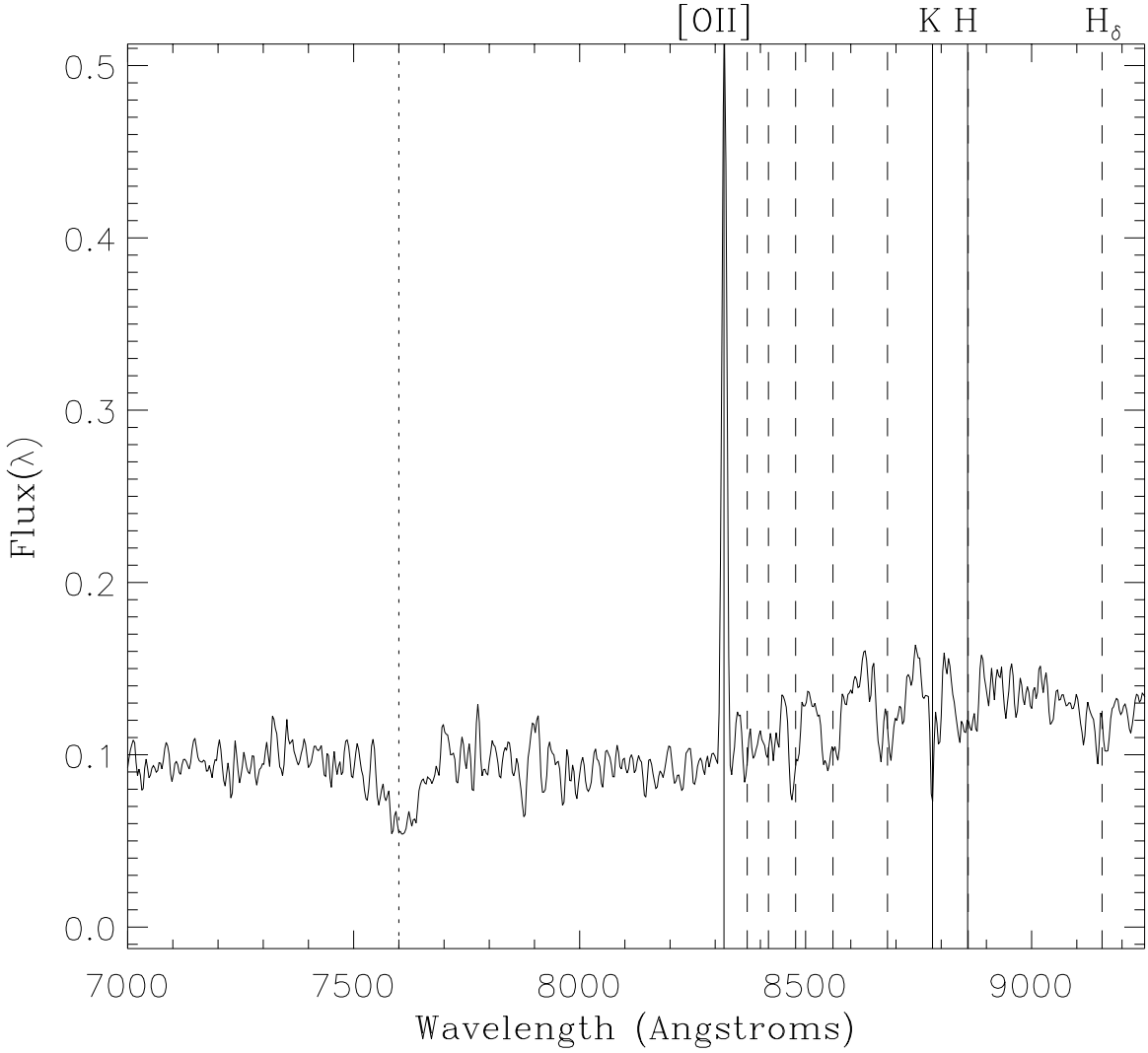

FIG. 12.— Co-added spectrum of the 17 spectroscpically confirmed star-forming members in RDCS J1252.9-2927. The most prominent spectral features in the displayed wavelength range are indicated. From right to left, the dashed lines are features of the Balmer series: H_δ, H_ϵ (next to the $CaII$ H line), and the higher order features H6, H7, H8, H9 and H10. The solid lines are the [OII] (λ3727) emission feature and the $CaII$ K and $CaII$ H absorption features. The dotted line indicates the A band telluric feature at 7600 Å. The spectrum has been smoothed by 5 pixels (1 pix $\sim$ 3.1Å).

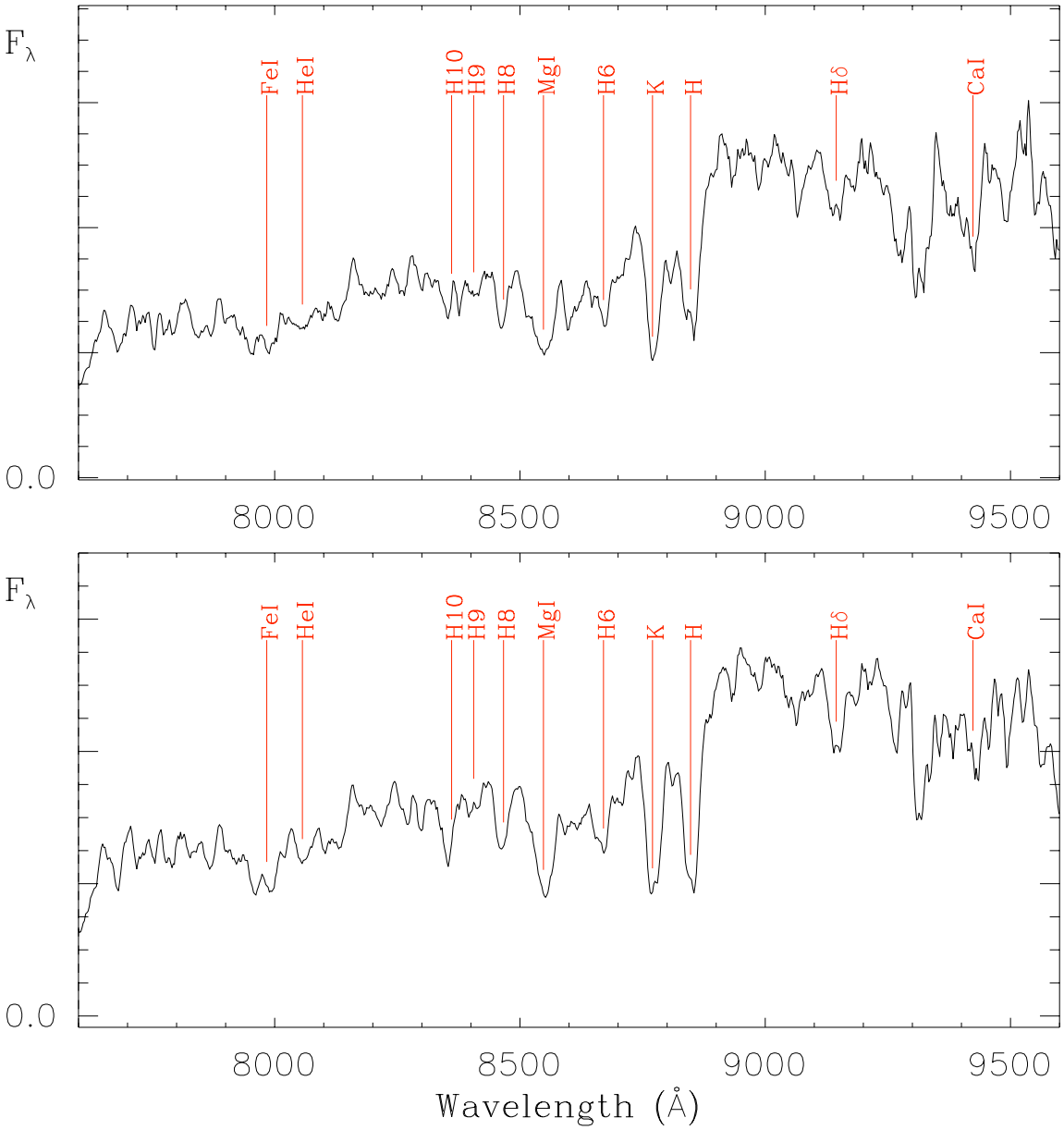

FIG. 13.— Co-added spectrum of the 10 (top) and 20 (bottom) brightest (in K_s) passive cluster members. Prominent Balmer absorption features become visible by co-adding spectra, indicating the existence of post-starburst stellar populations in passive early-type galaxies at this redshift (see footnote in table 2 for a list of spectral features). The spectra have been smoothed by 5 pixels (1 pix $\sim$ 2.5Å).

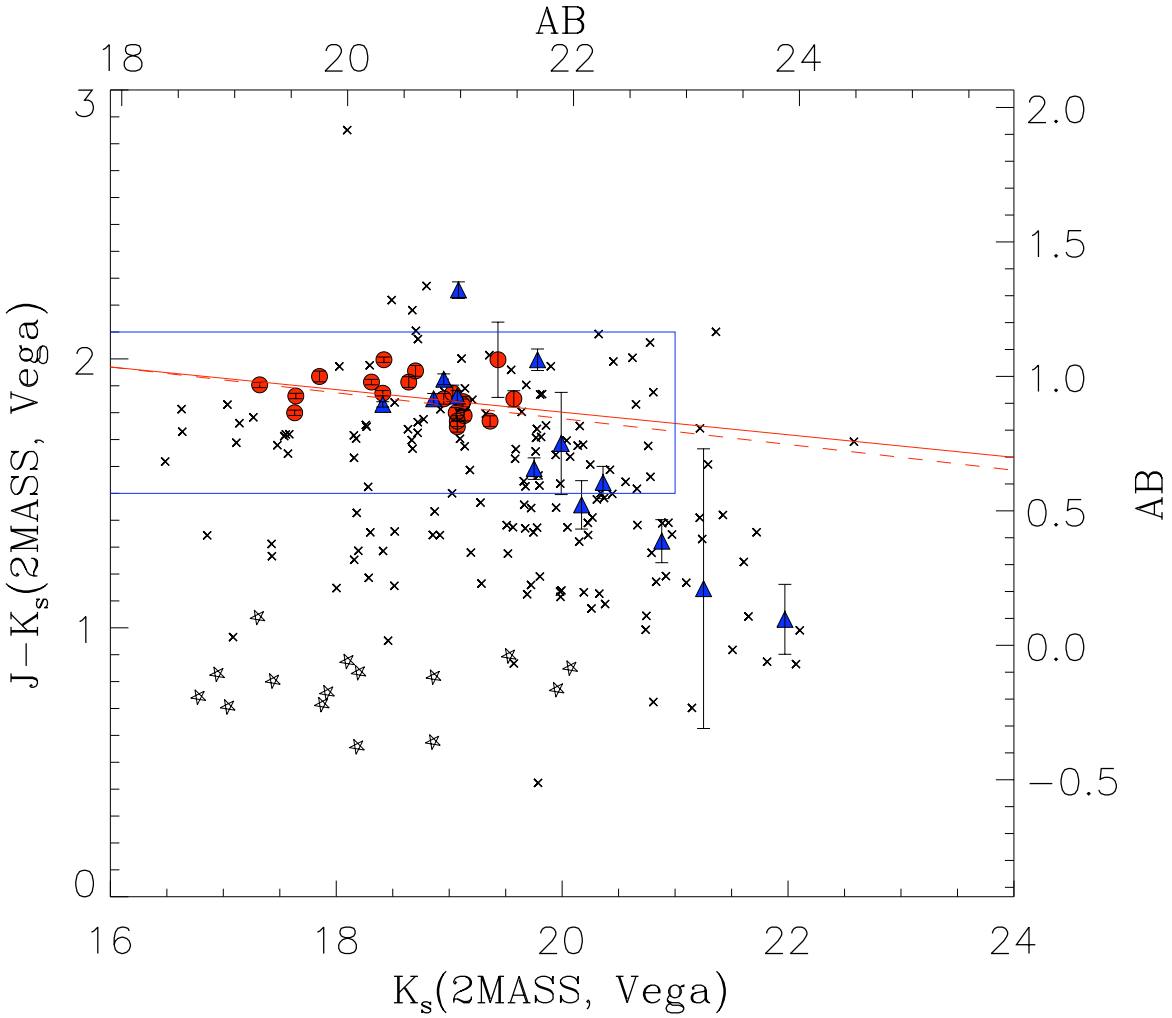

FIG. 14.— Near-IR (reprocessed ISAAC data) color-magnitude diagram (CMD) of spectroscopic cluster members. Colors and magnitudes are shown in the 2MASS (Carpenter 2001; Cohen et al. 2003) and AB systems. The star symbols represent spectroscopically confirmed stars. Crosses are non-cluster members, i.e., objects with redshift $z \leq 1.22$ or $z \geq 1.25$. Filled red circles are spectroscopic cluster members (objects with redshift in the range $1.22 < z < 1.25$) without detectable [OII](λ3727) emission. The filled blue triangles are cluster members with [OII] emission. The dotted red line is the fit published in Lidman et al. (2004) including only galaxies within 20″ (0.17 Mpc) of the cluster center and within the blue rectangle. The solid red line (fit #2 in Table 6) is the fit presented in this work using only cluster members without [OII] (filled red circles). The slope and scatter about the two fits are listed in Table 6. The reddest star-forming member in the diagram corresponds to the confirmed AGN (ID=174). In this study we extend the analysis of the color-magnitude diagram of RDCS J1252.9-2927 to more than 1 Mpc in radius from the cluster center, and show that the red sequence extends to more than 0.5 Mpc in clustercentric radius.

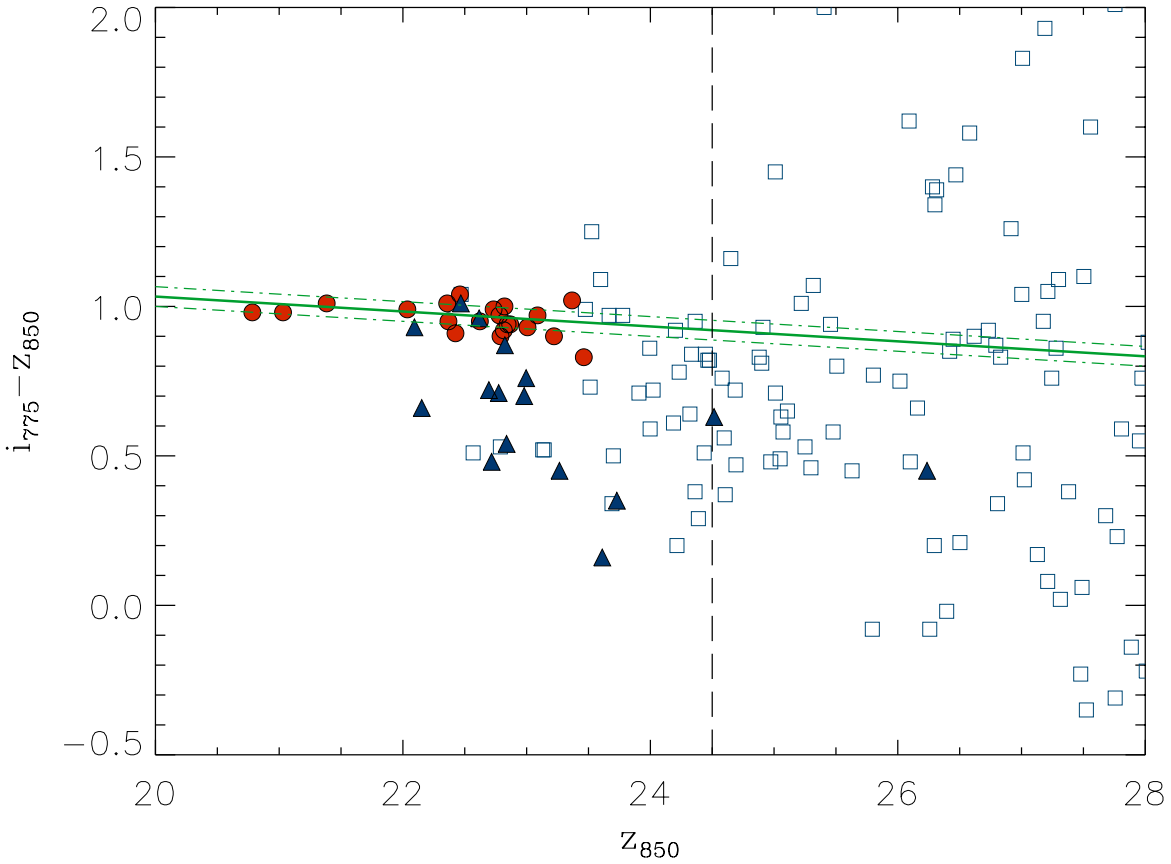

FIG. 15.— ACS color-magnitude diagram (CMD) of spectroscopic cluster members (circles and triangles) and photometric members (squares). Photometric redshifts are computed by using the Bayesian method (Benítez 2000; Benítez et al. 2004). Colors and magnitudes are computed in the AB system. Red circles correspond to passive members while blue triangles correspond to [OII] emission line galaxies. The best-fit color-magnitude relation and scatter from Blakeslee et al. (2003b) are shown in green. Only objects within the central 1 Mpc region around the cluster center and with $1.14 < z_{phot} < 1.34$ are shown. A color-magnitude relation for early-type galaxies is seen down to $z_{850} = 24.5$ (dashed vertical line). At fainter magnitudes this relation seems to be truncated in the bandpasses shown here, although this cutoff magnitude can be affected by uncertainties in the photometric redshifts.

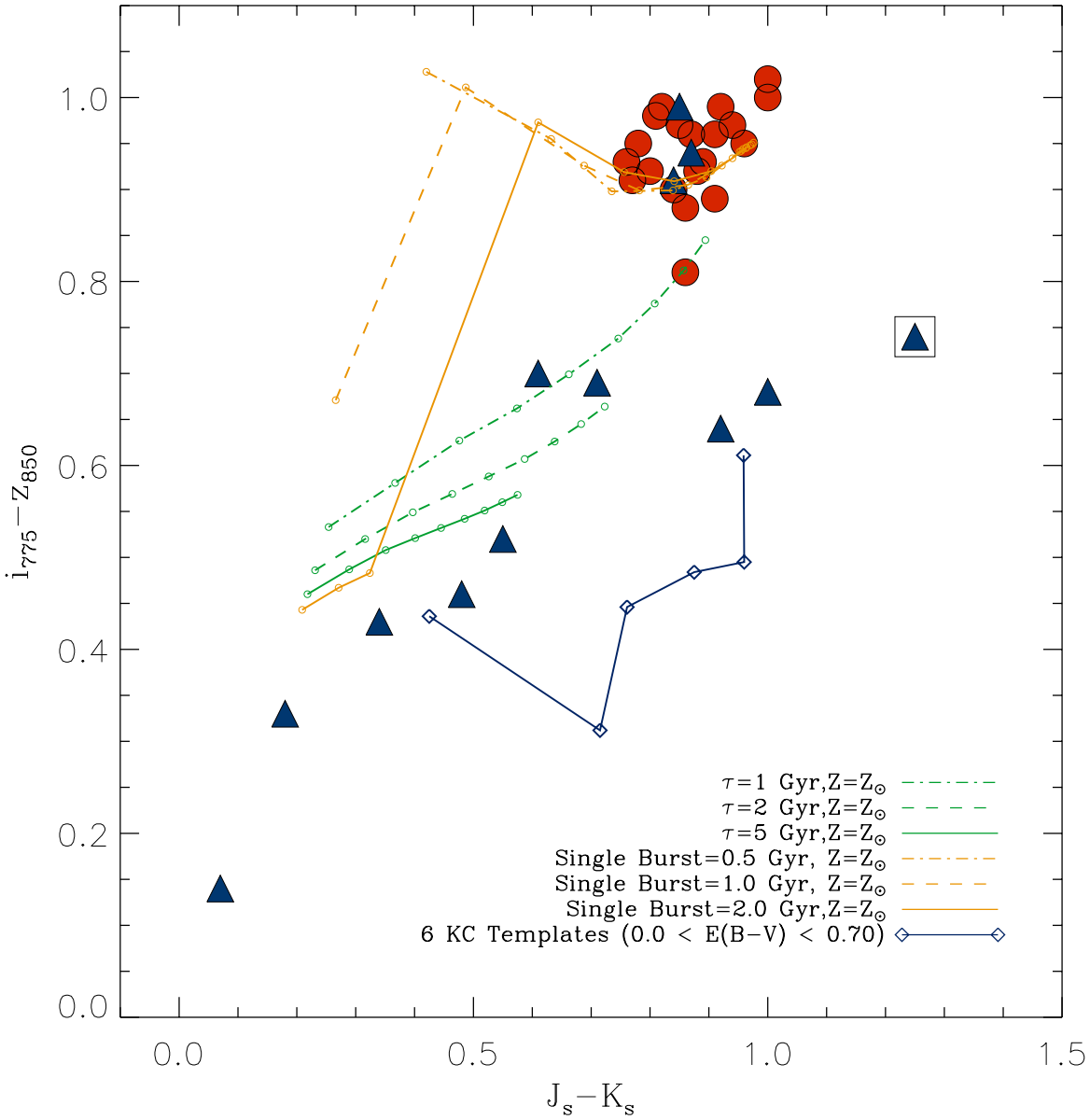

FIG. 16.— Distribution of spectroscopic cluster members in color-color space. All magnitudes are AB. Red filled circles correspond to passive galaxies and blue filled triangles to [OII] emission galaxies. The open square indicates the only known AGN at the cluster redshift (ID=174). Model tracks are the same as those used in Fig. 2, spanning from 1 Gyr to 5 Gyr with intervals of 0.5 Gyr (see text). In this plot we show at the same time the color of passive and star-forming galaxies in both ACS and ISAAC filters. We note that this color-color diagram improves the selection of passive, red objects with respect to a $J - K$ color cut only.